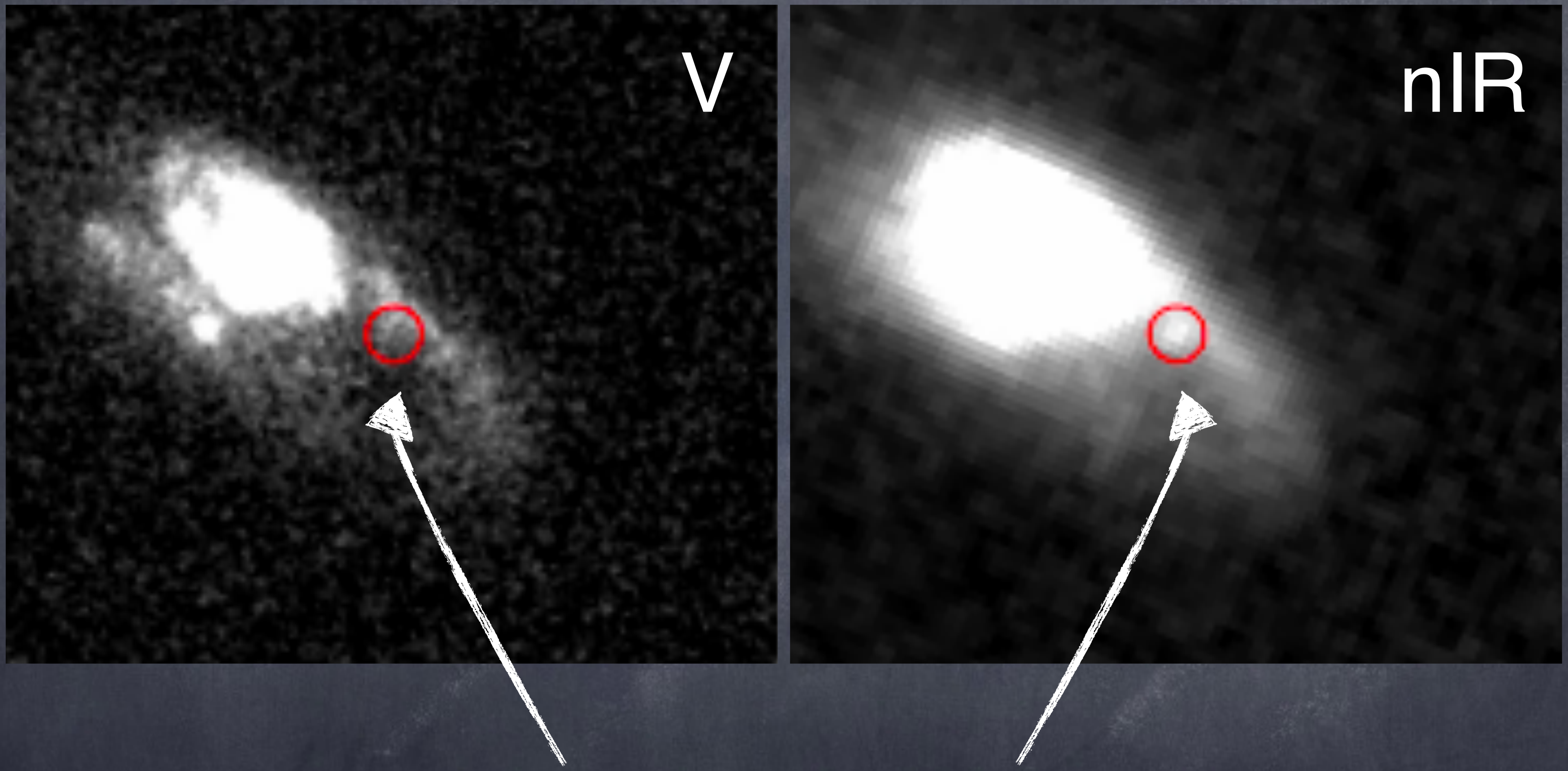
Mergers, Gamma-Ray Bursts and Gold

Tsvi Piran

The Hebrew University

David Eichler, Mario Livio, David Schramm, Doron Grossmam,
Stephan Rosswog, Oleg Korobkin, Ehud Nakar,
David Wanderman Ben Margalit

The Hubble Space Telescope
June 13th 2013

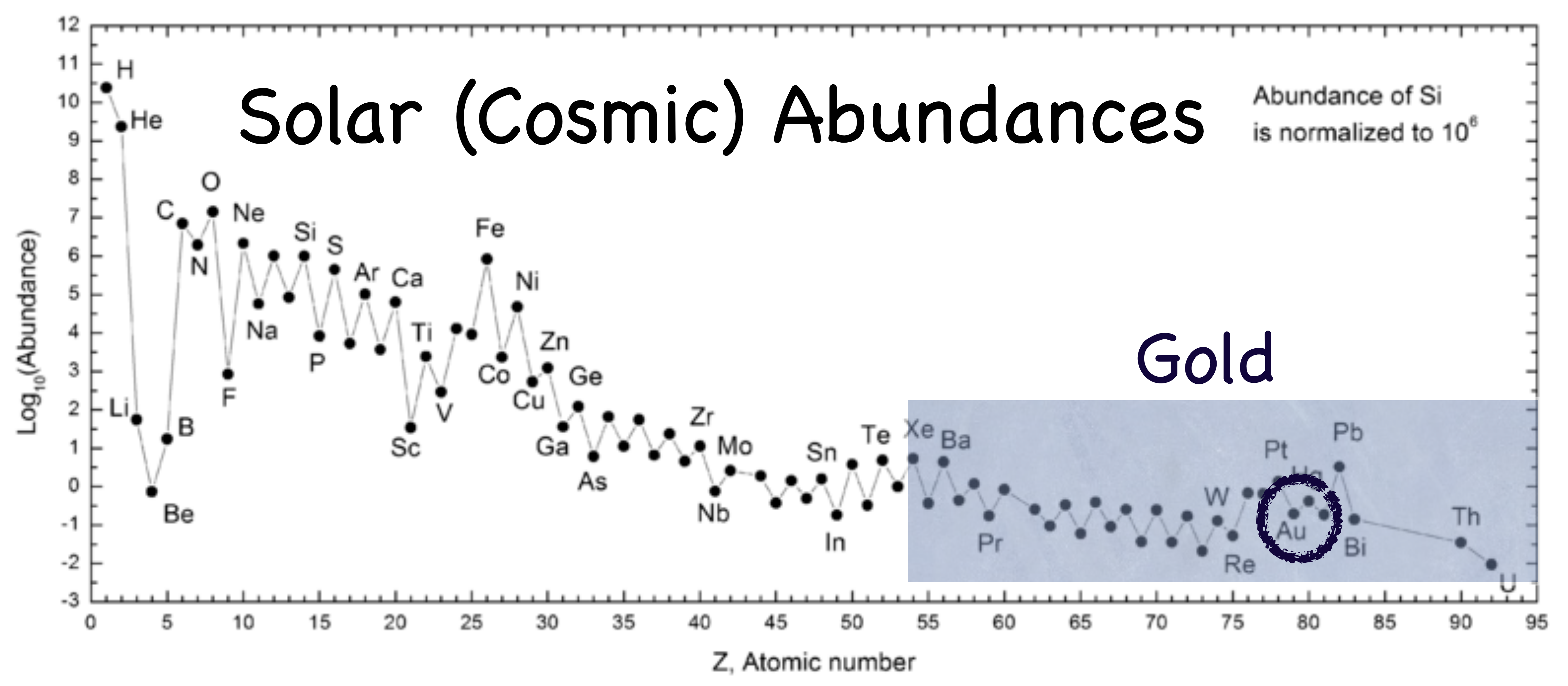


Is this the “smoking gun” proving the origin of Gold (and other heavy elemets) in the Universe?

Outline

1. Nucleosynthesis 101
2. Neutron Stars and Mergers
3. Gamma-Ray Bursts
4. The Li-Paczynski Macronova (kilonova)
5. Putting it all togather - GRB 130603B
6. The origin of Gold

1. Nucleosynthesis 101



Solar (Cosmic) Abundances

Gold

How are these elements produces?

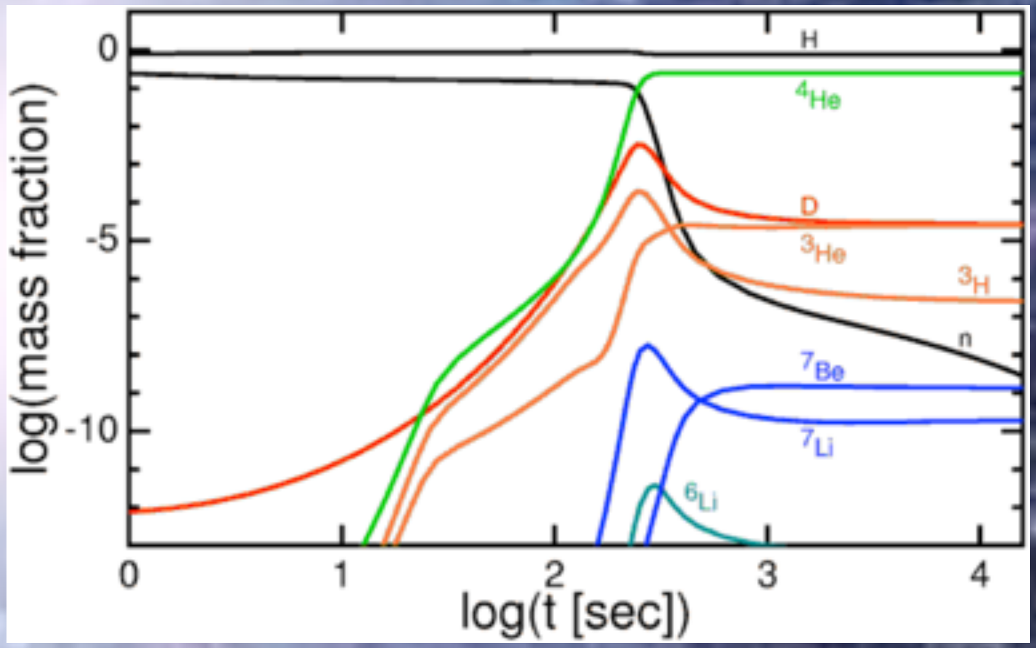
BB (Big Bang) Nucleosynthesis

- 24% of the Universe is He.
- This He is produces in the big Bang. ✓

George Gammow

Only* He is produced in the big bang

*and minute quantities of ^{2}H, ^{3}He, ^{3}H, Li and Be

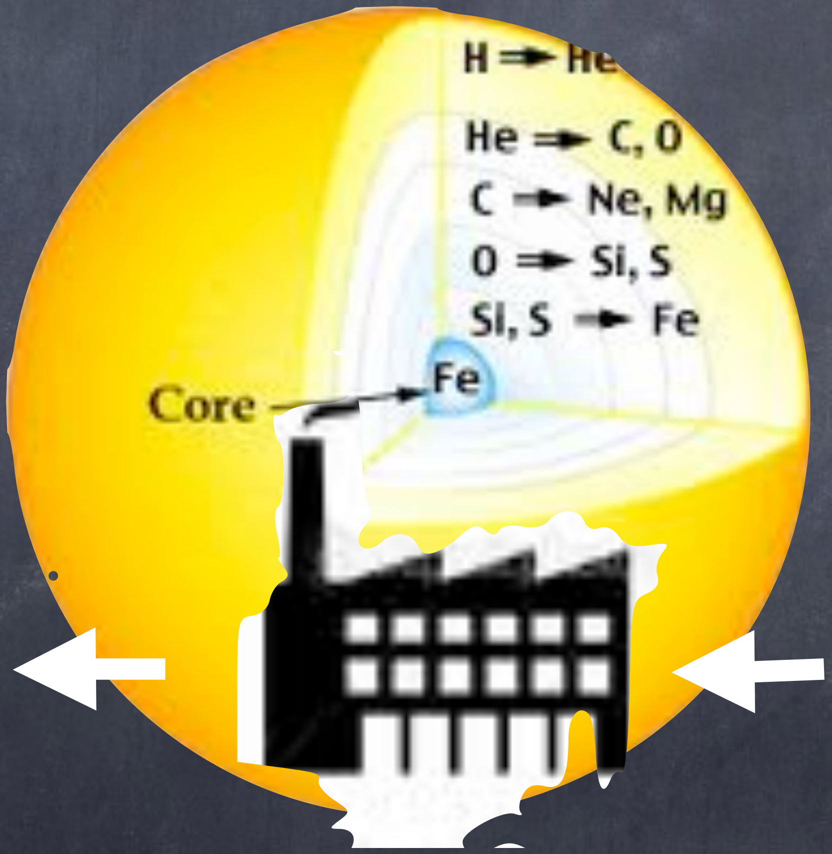


Peebles, Adouze, Schramm, Steigman...

How did other elements form?

Burbidge, Burbidge, Fowler and Hoyle

B^2FH 1957



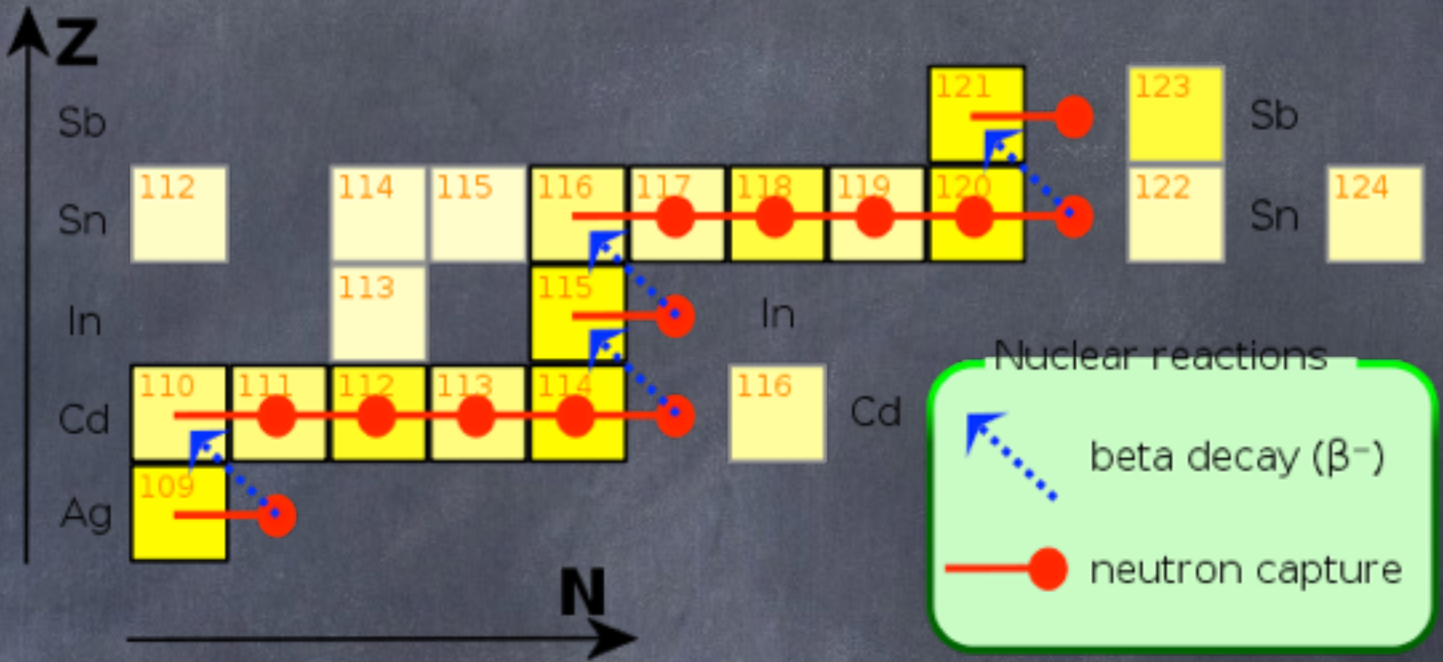
He,C,O,Ne,Mg
Si,S,Fe,Ni.....

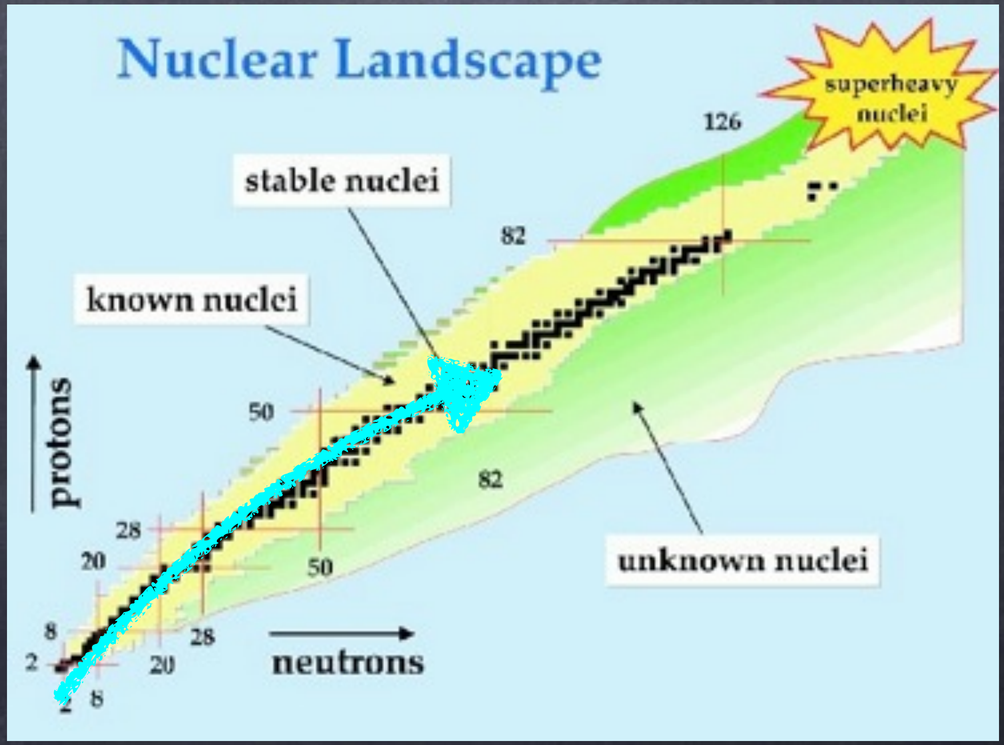
H

Elements up to Iron are produced in stars

S (slow) Process

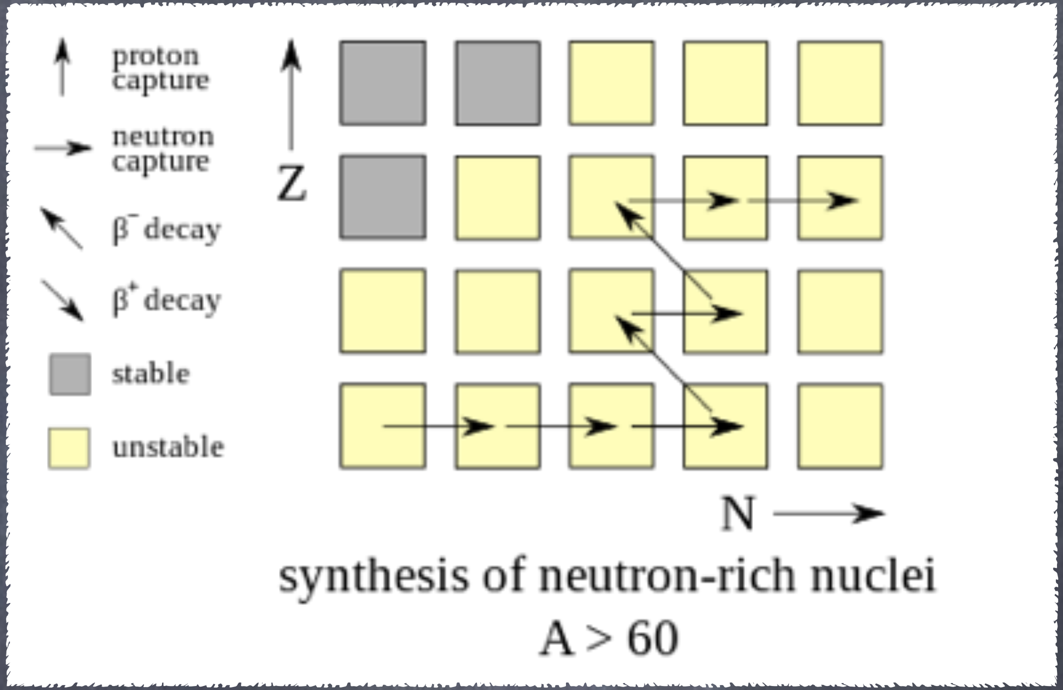
- Neutron capture slower than beta decay.
- Low neutron densities.
- time scale - years.
- Moves along the valley of nuclear stability.
- Final abundances depend on the conditions within the site.



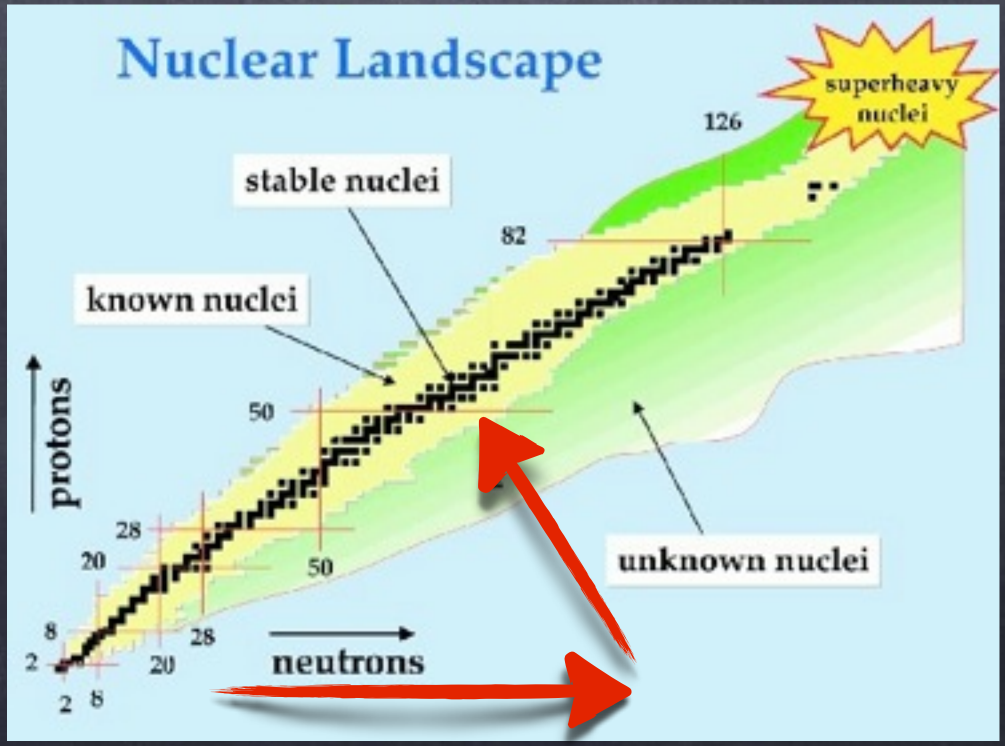


r (rapid) Process

- Neutron capture faster than beta decay.
- High neutron densities.
- Time scales - seconds.
- On the neutron rich side of nuclear stability.
- Uniform final abundances.

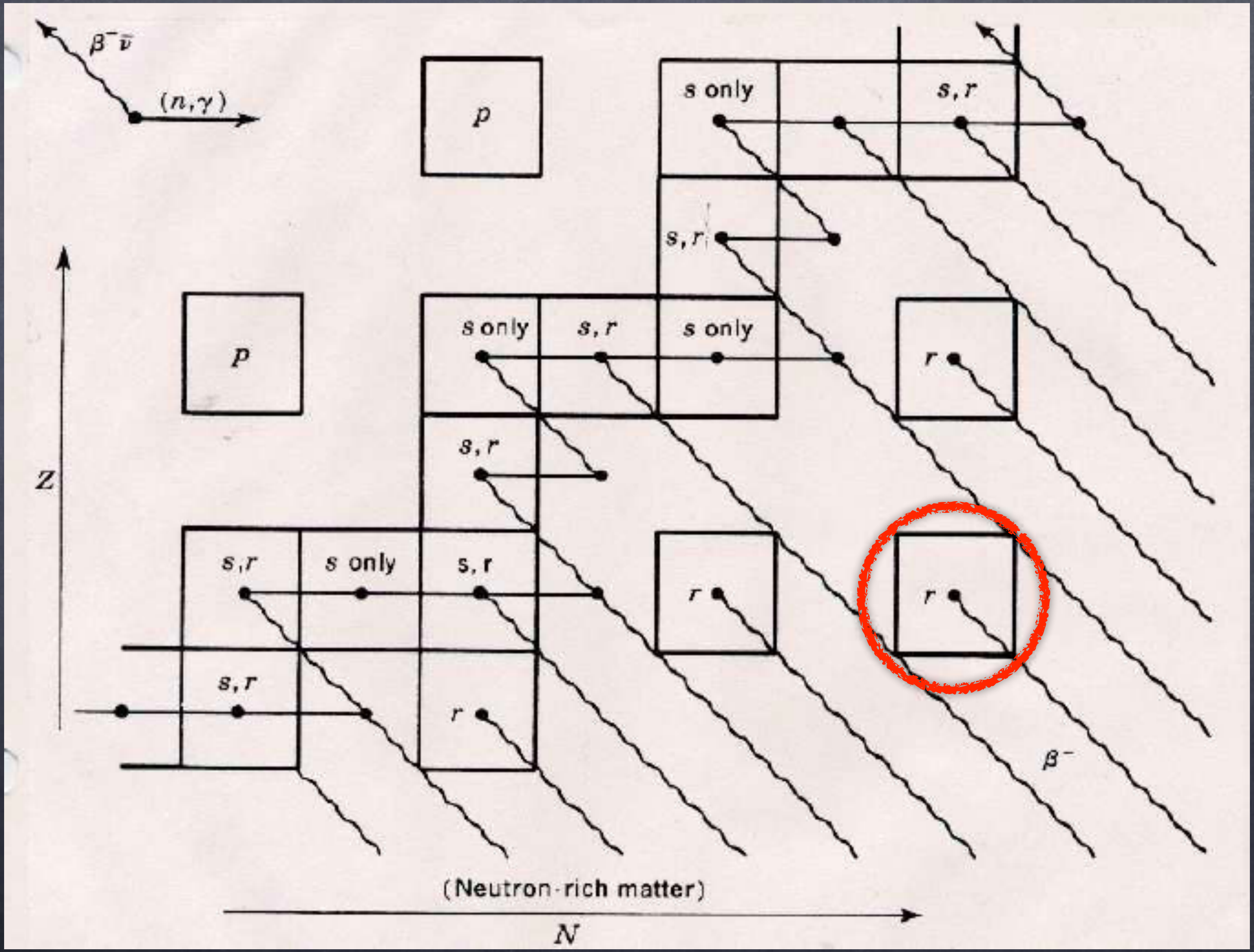


synthesis of neutron-rich nuclei
$A > 60$



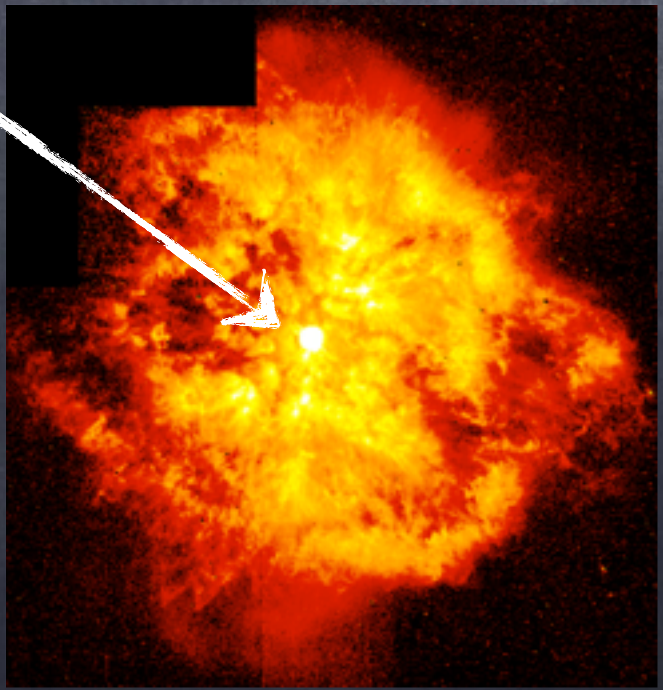
Nuclear Landscape

s and r processes

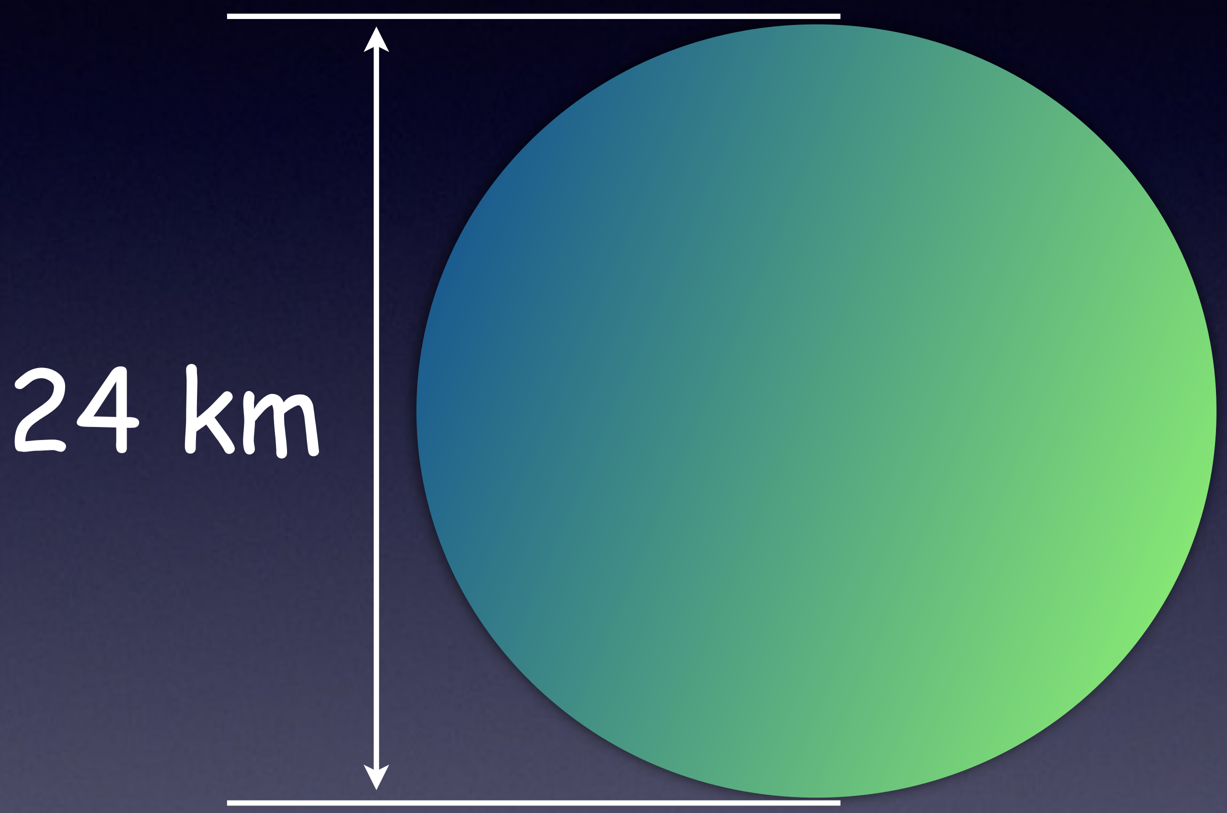


Explosive r-process

- ν flux from the newborn neutron star produce excess of neutrons in **Supernova explosion.**



2. Neutron stars and mergers

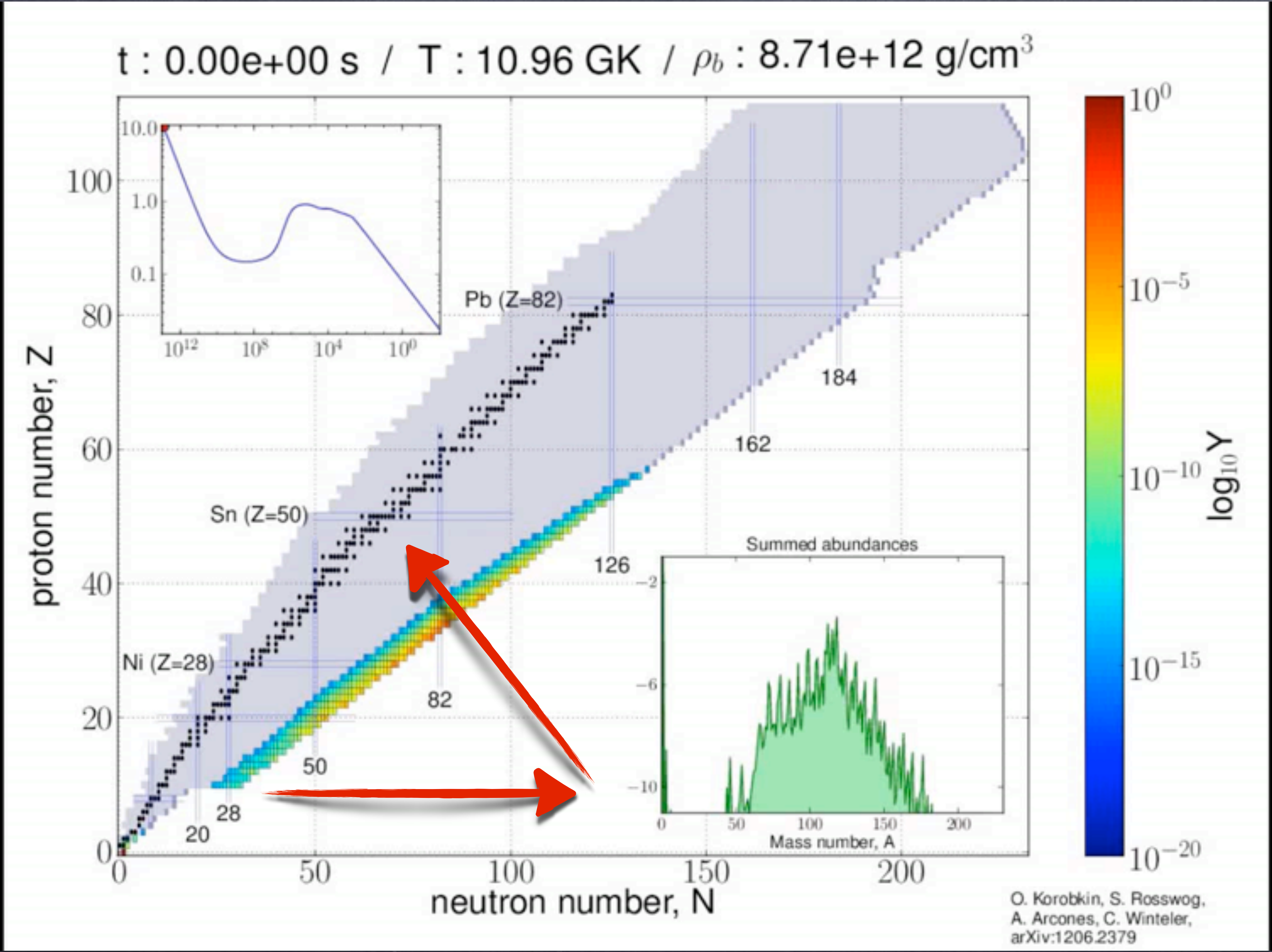


95% neutrons!

1 cc of neutron star material

$n \rightarrow p + e + \nu$

Decay of neutron star matter

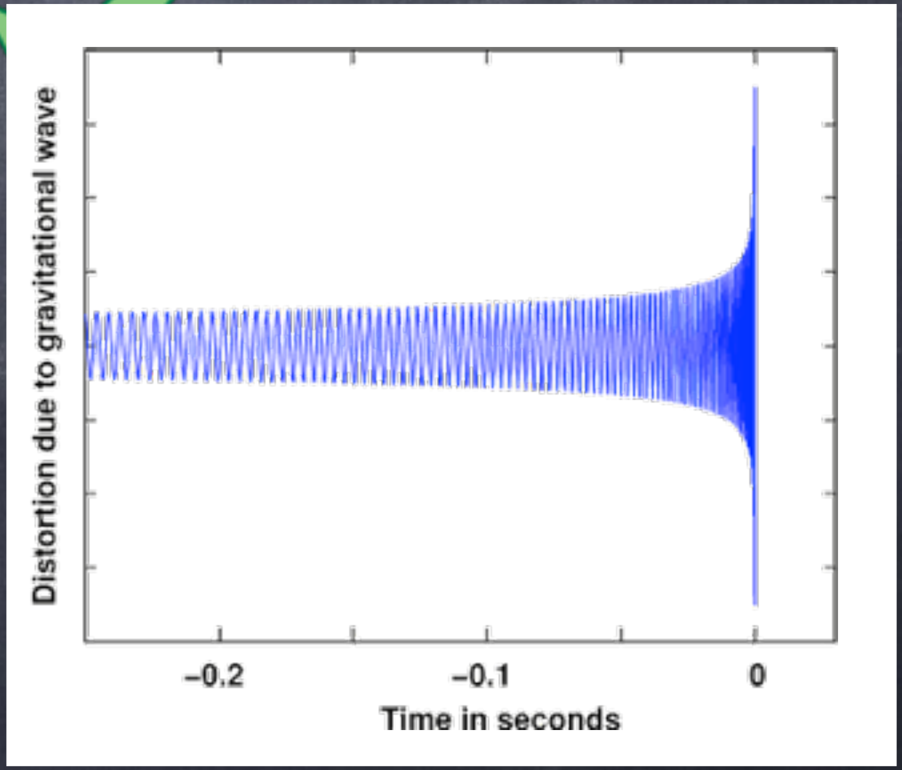


O. Korobkin, S. Rosswog, A. Arcones, C. Winteler, arXiv:1206.2379

Neutron Star Mergers

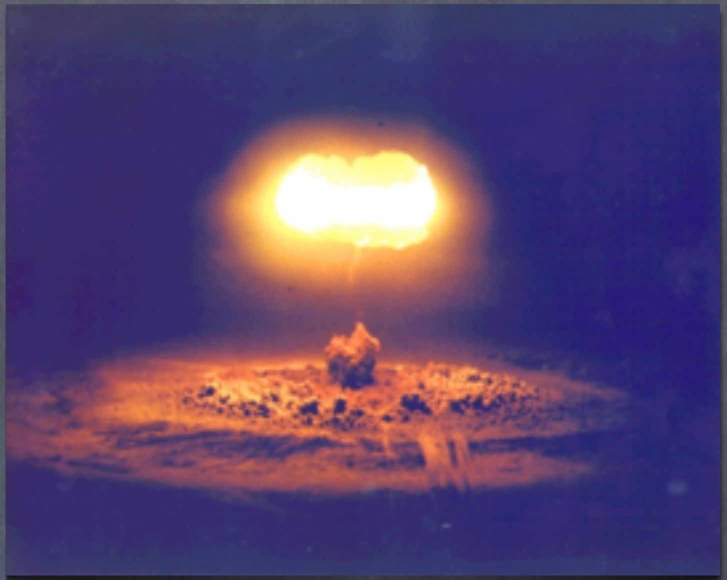
Binary Neutron Stars

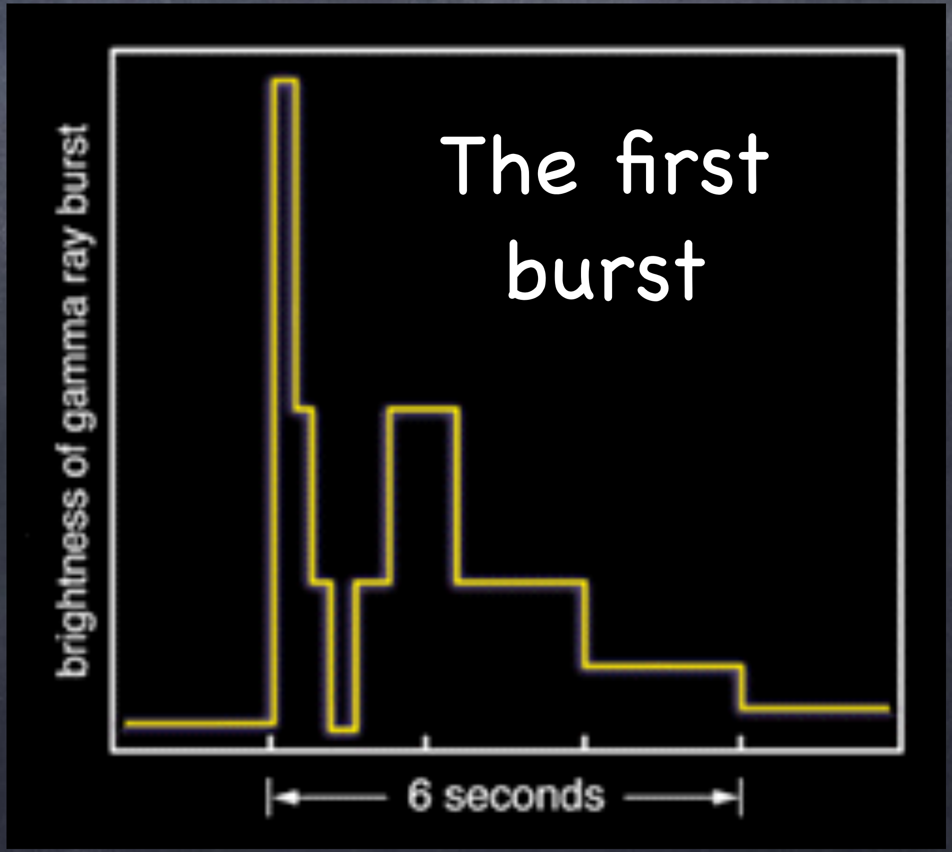
$$\frac{dr}{dt} = -\frac{64}{5}\frac{G^3}{c^5}\frac{(m_1 m_2)(m_1 + m_2)}{r^3}$$



3. Gamma Ray Bursts

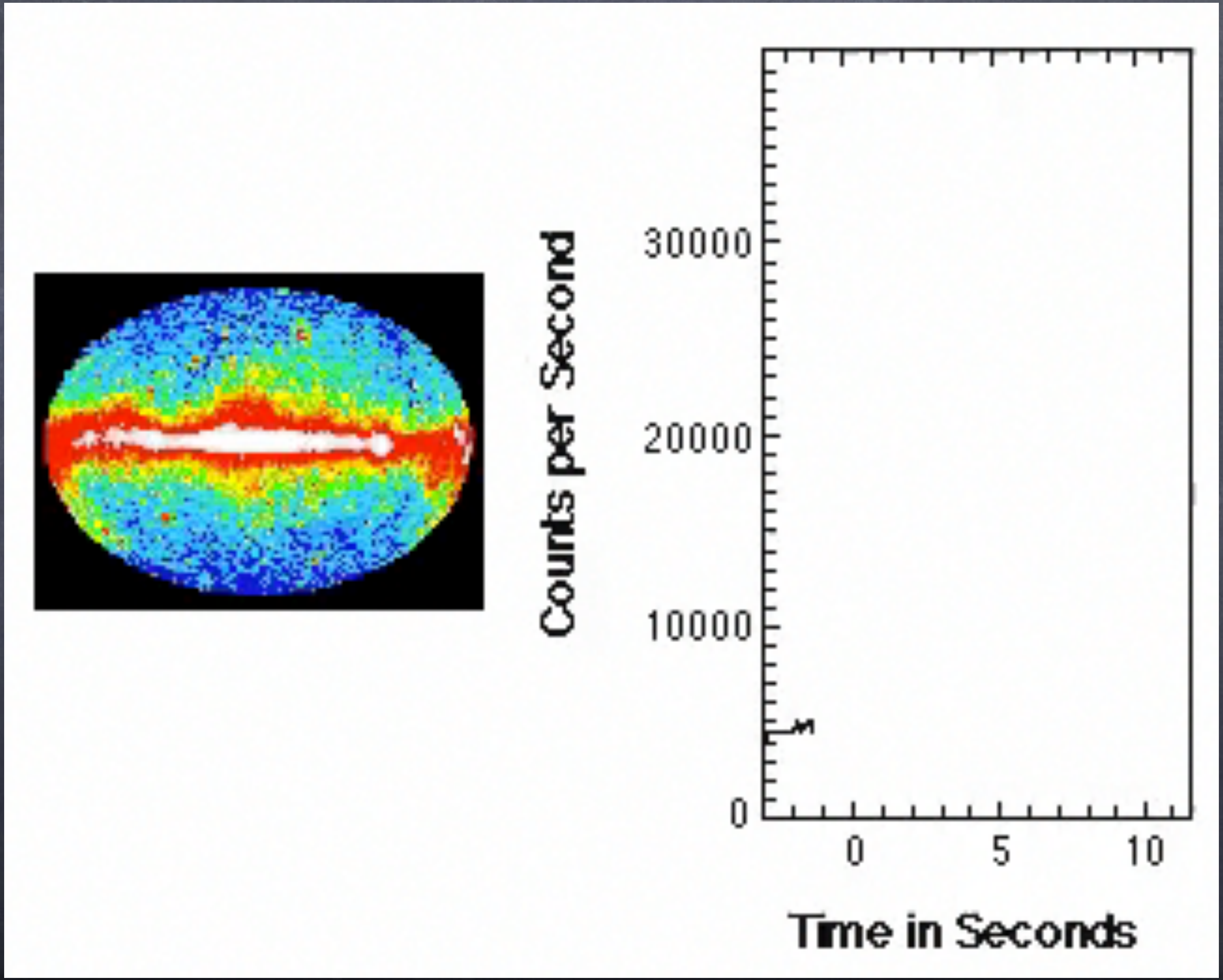
The Vela Satellites



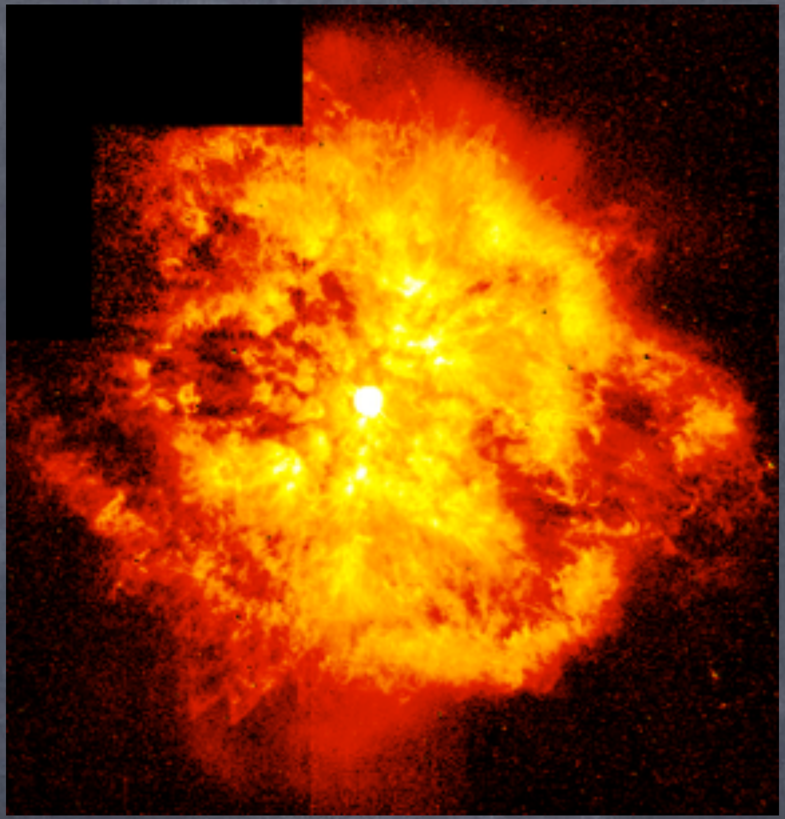


The first burst

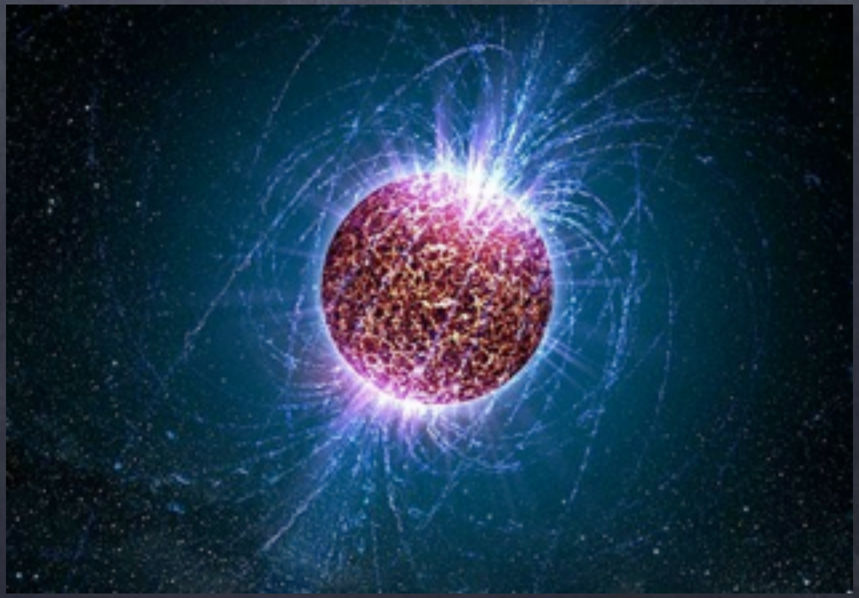
The sky in gamma-Rays



The late 80ies

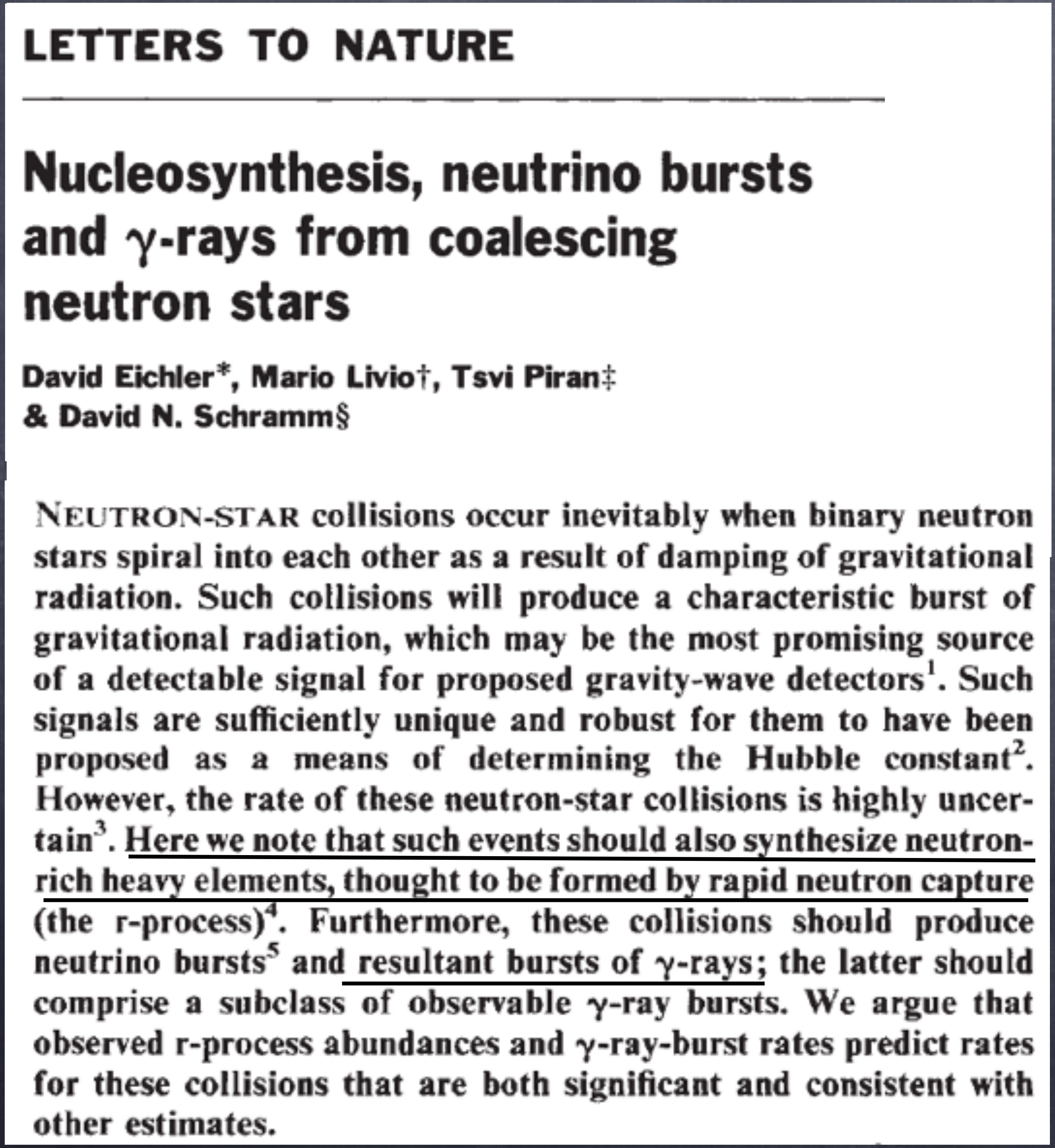


- r-process material from Supernovae



- GRBs from magnetic flares on galactic neutron stars (E~10^{40} ergs).

Two provocative ideas



LETTERS TO NATURE

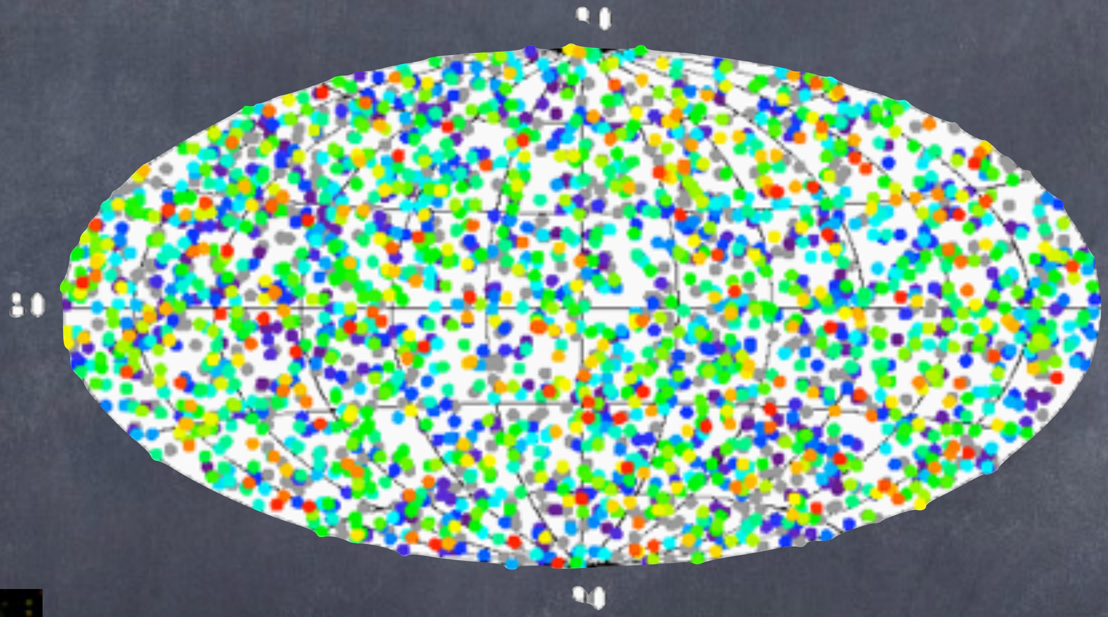
Nucleosynthesis, neutrino bursts and γ-rays from coalescing neutron stars

David Eichler*, Mario Livio†, Tsvi Piran‡
& David N. Schramm§

NEUTRON-STAR collisions occur inevitably when binary neutron stars spiral into each other as a result of damping of gravitational radiation. Such collisions will produce a characteristic burst of gravitational radiation, which may be the most promising source of a detectable signal for proposed gravity-wave detectors[1]. Such signals are sufficiently unique and robust for them to have been proposed as a means of determining the Hubble constant[2]. However, the rate of these neutron-star collisions is highly uncertain[3]. Here we note that such events should also synthesize neutron-rich heavy elements, thought to be formed by rapid neutron capture (the r-process)[4]. Furthermore, these collisions should produce neutrino bursts[5] and resultant bursts of γ-rays; the latter should comprise a subclass of observable γ-ray bursts. We argue that observed r-process abundances and γ-ray-burst rates predict rates for these collisions that are both significant and consistent with other estimates.

90ies: GRBs are cosmological

1992: BATSE - GRBs have a coslomogical distribution



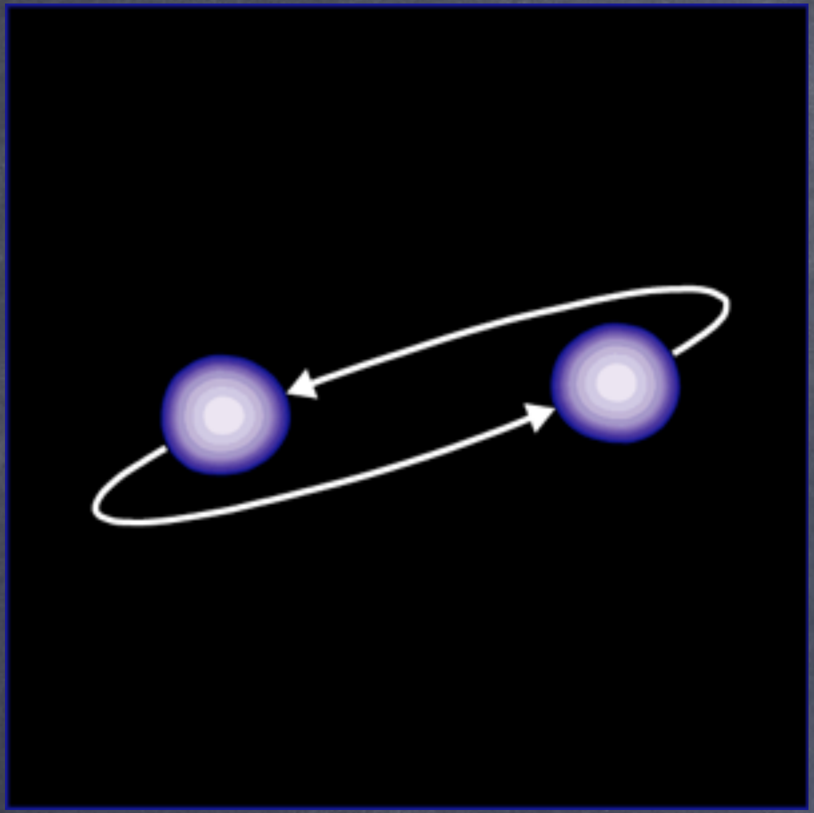
Gamma-Ray Bursts

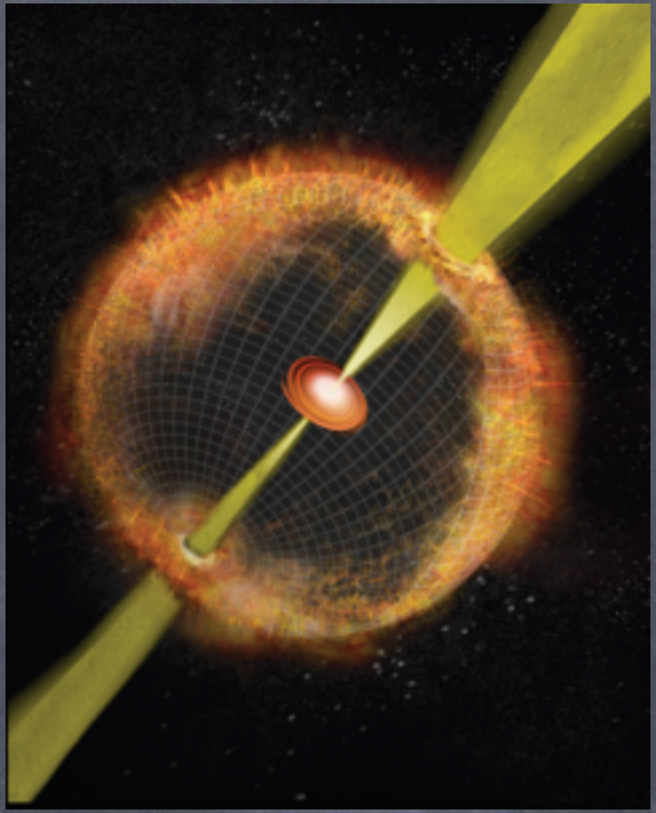
1997: BeppoSAX - GRBs' afterglow that enables redshift measurements confirming the coslomogical origin

1988	2013
~~r-process from Supernovae~~	Supernovae cannot produce A>130
~~GRBs from magnetic flares on galactic neutron stars~~ (E ~10^{40} ergs).	GRBs are cosmological (E ~10^{51} ergs).

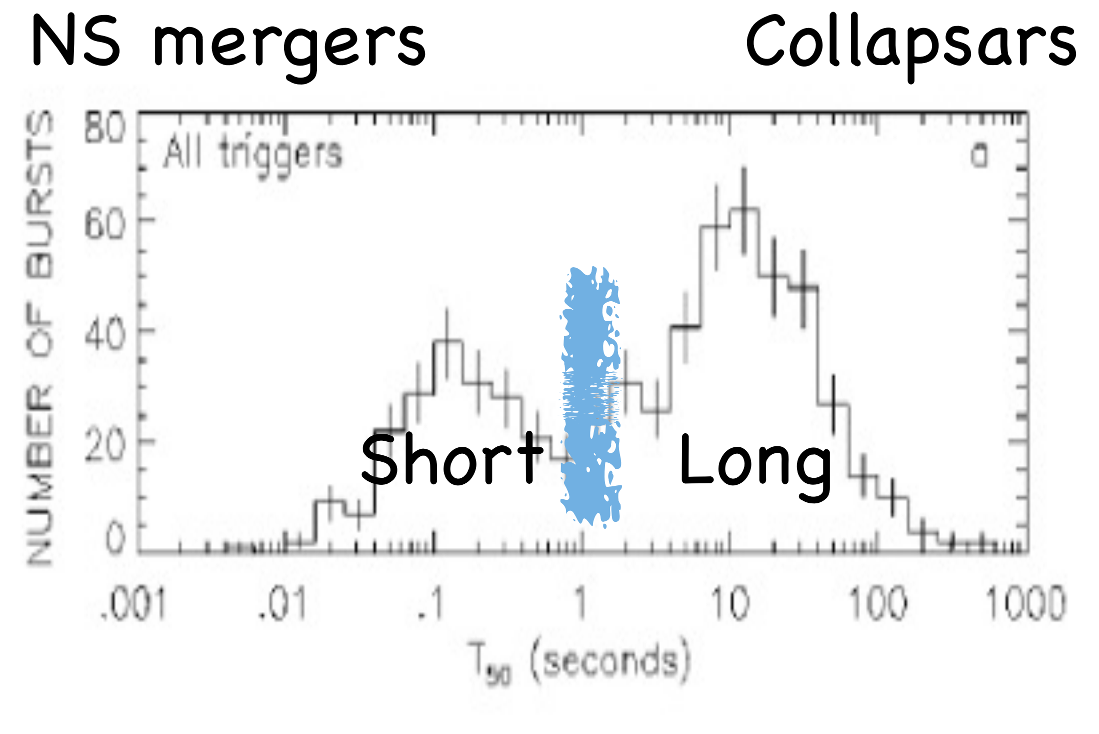
Eichler, Livio, TP,
Schramm, 88

MacFadyen & Woosley,
98



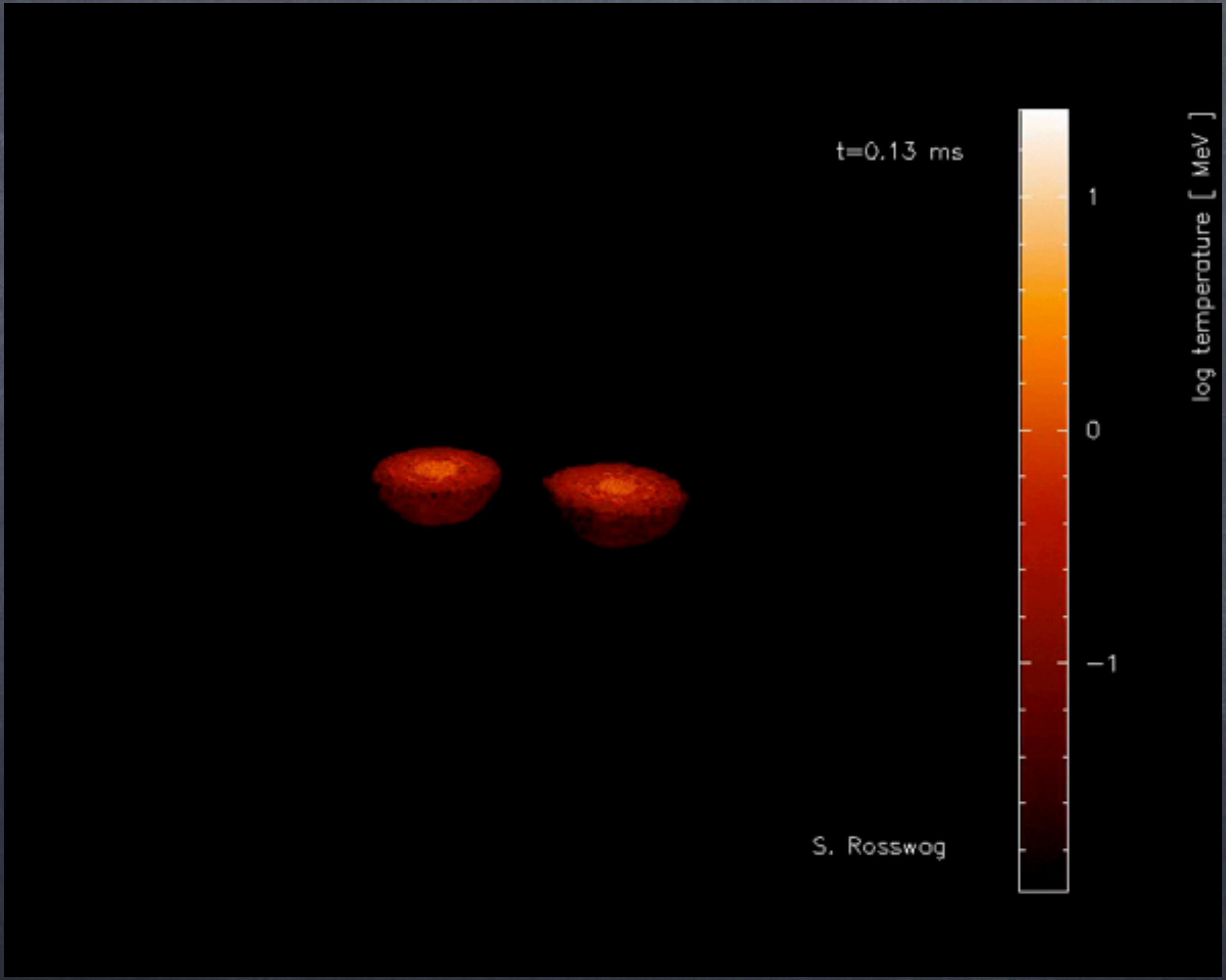


Indirect
Evidence



Direct
Evidence

Mergers ejects 0.01-0.04M$_{sun}$ with E$_{k}$ ~ 10^{50}-10^{51} ergs



Stephan Rosswog

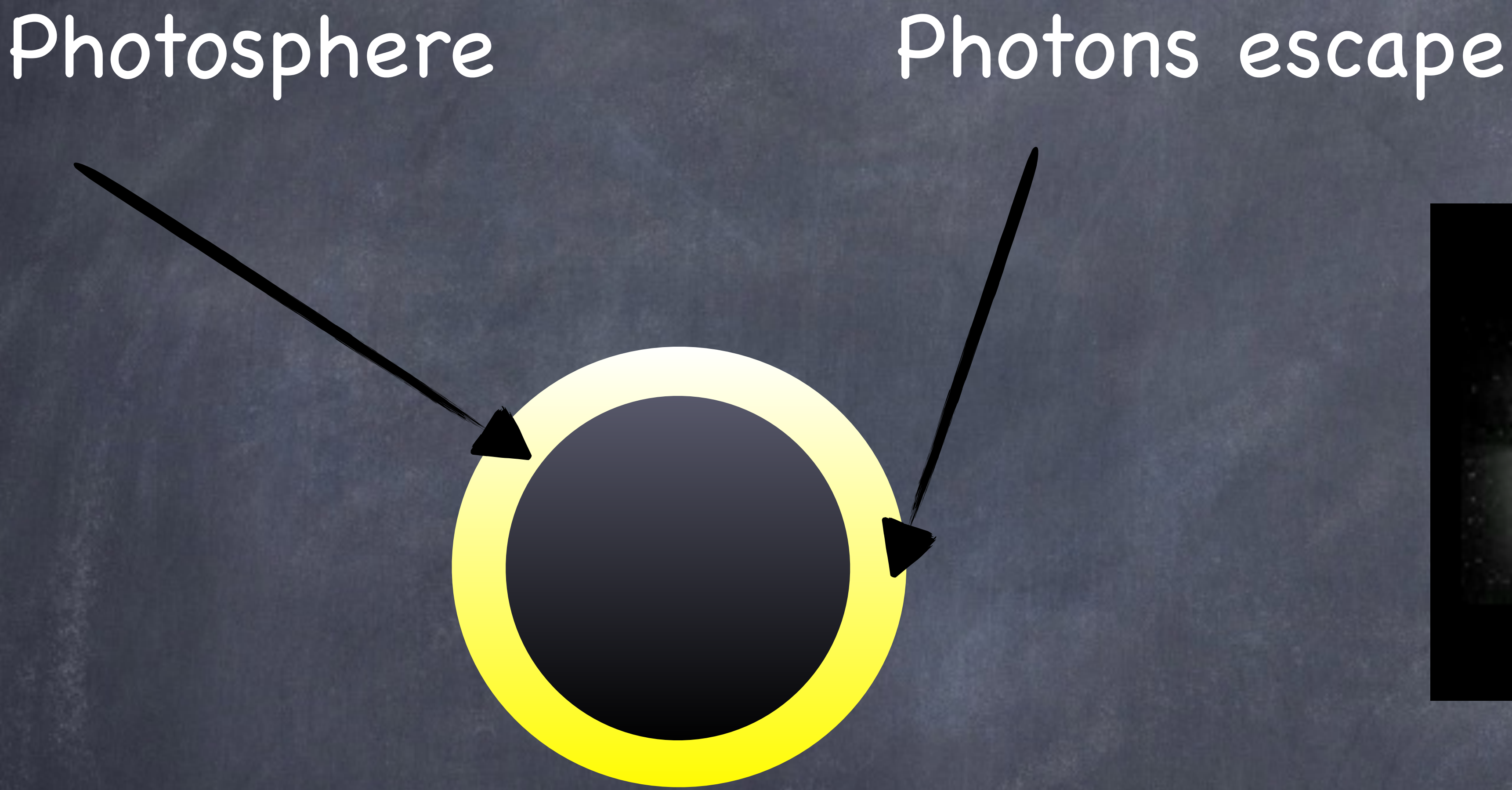
4. Macronova* (Li & Paczynski 1997)

Bohdan Paczynski

- Radioactive decay of the neutron rich matter.
- $E_{radioactive} \approx 0.001\ Mc^2 \approx 10^{50}$ erg
- A weak short Supernova like event.
- Macronovae follow short GRBs but could appear without a short GRB as those are beamed.

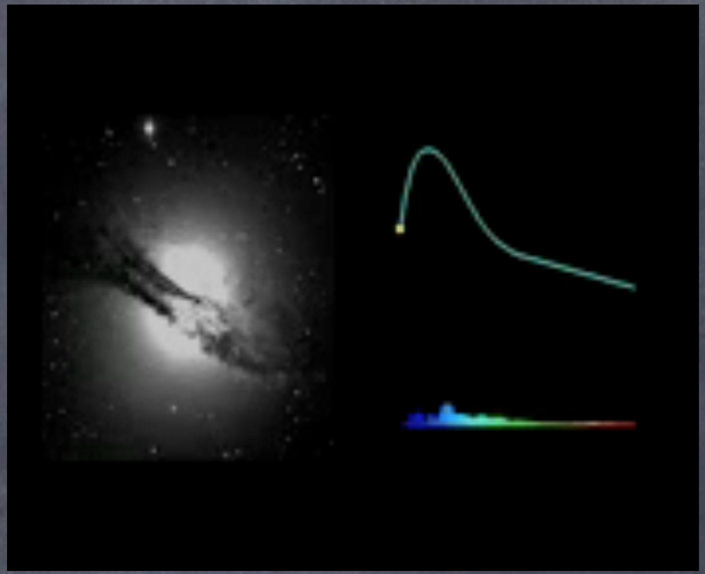
*Also called Kilonova

Supernova

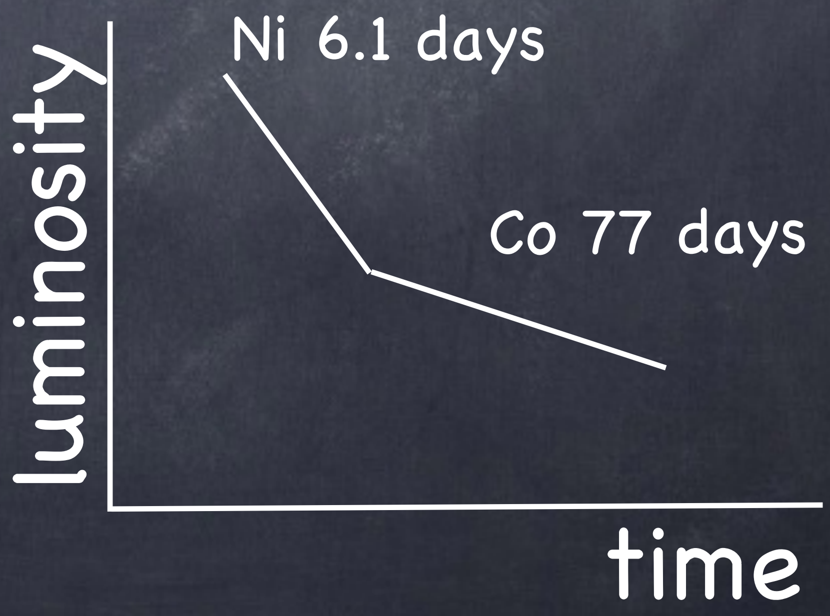


Photosphere

Photons escape

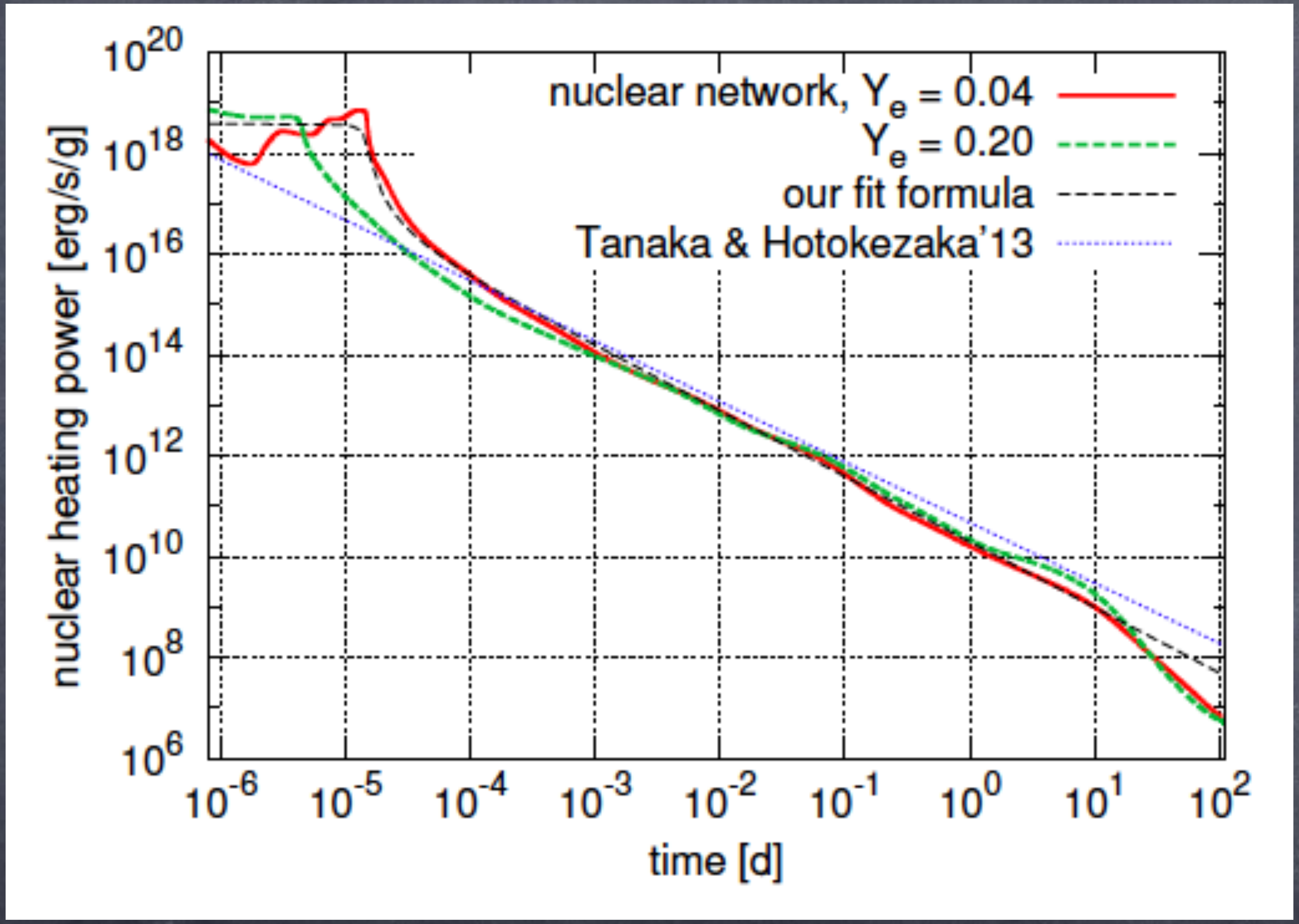


Powered by radioactive decay of ^{56}Ni->^{56}Co->^{56}Fe

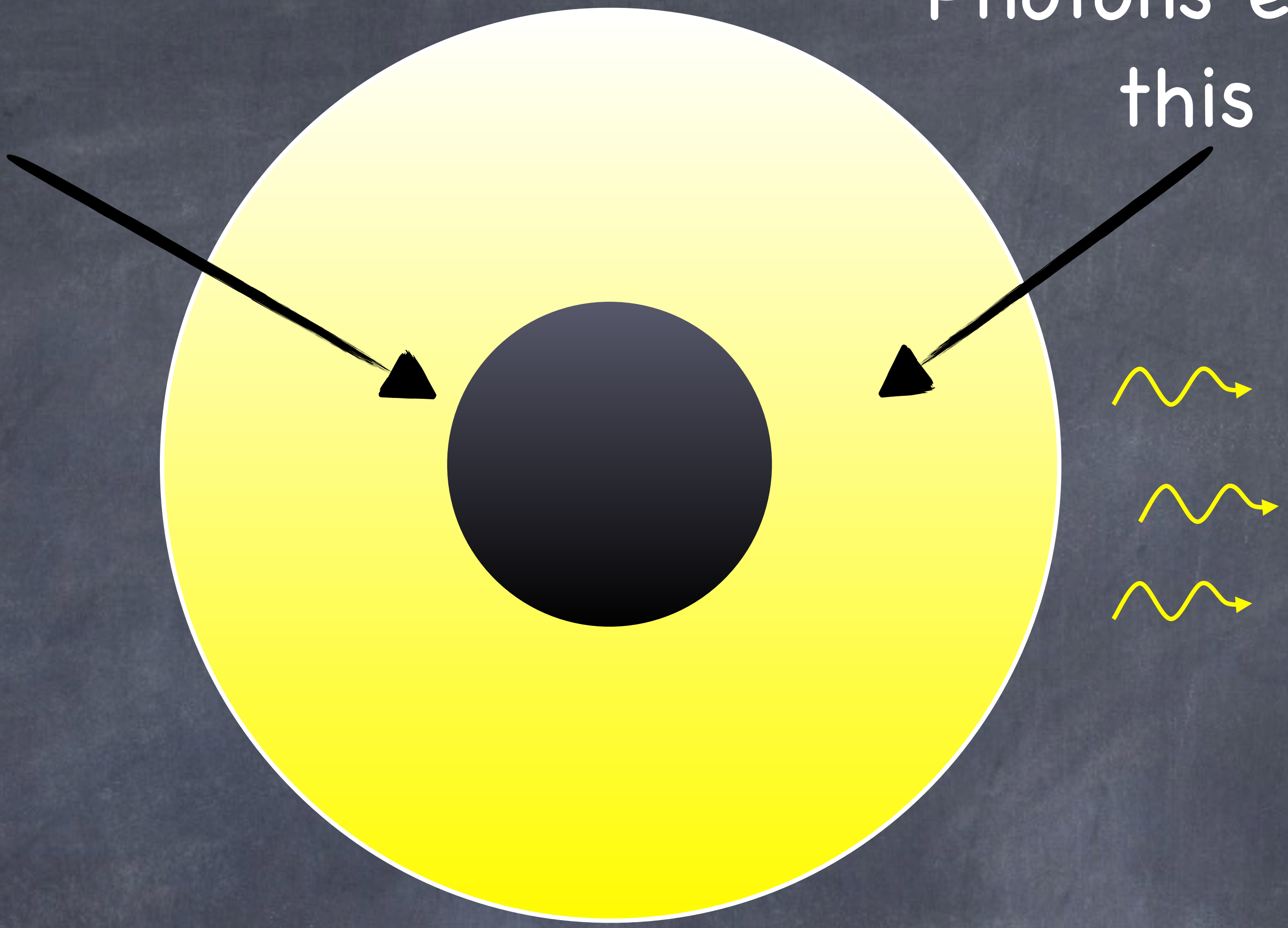


Radioactive Decay

Korobkin + 13; Rosswog, Korobkin + 13

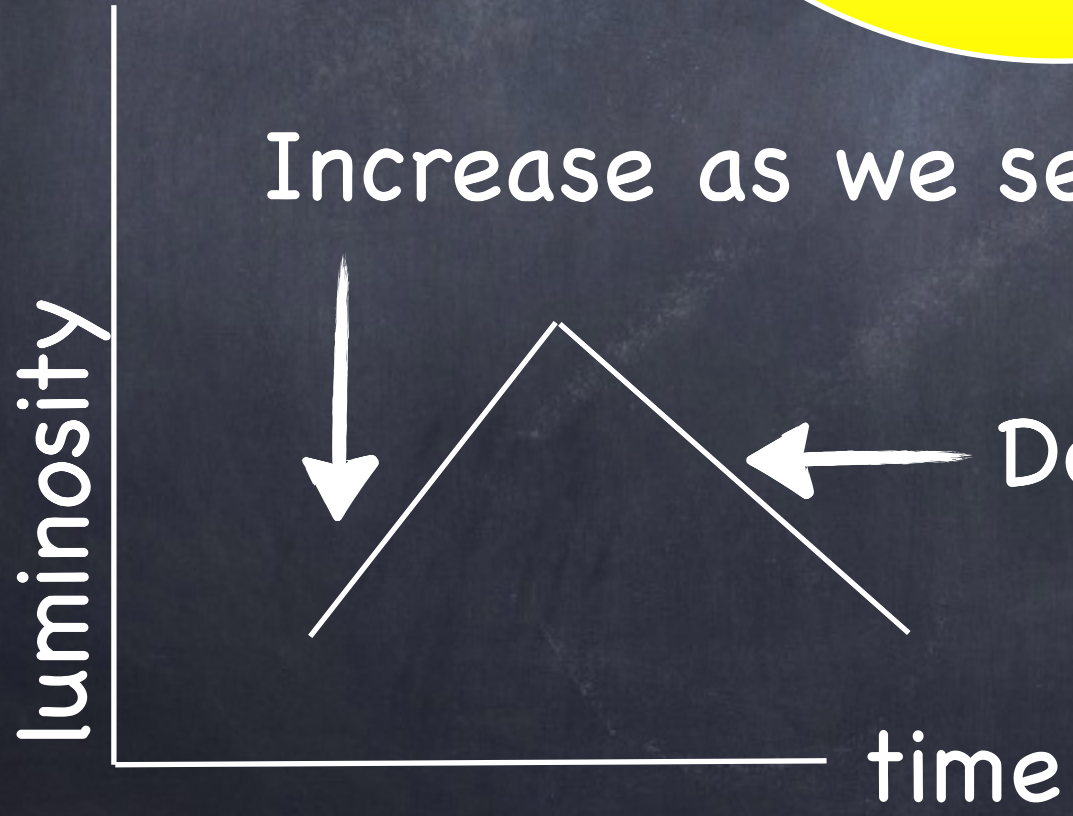


- After a second $dE/dt \propto t^{-1.3}$ (Freiburghaus+ 1999; Korobkin + 2013)



τ=c/v

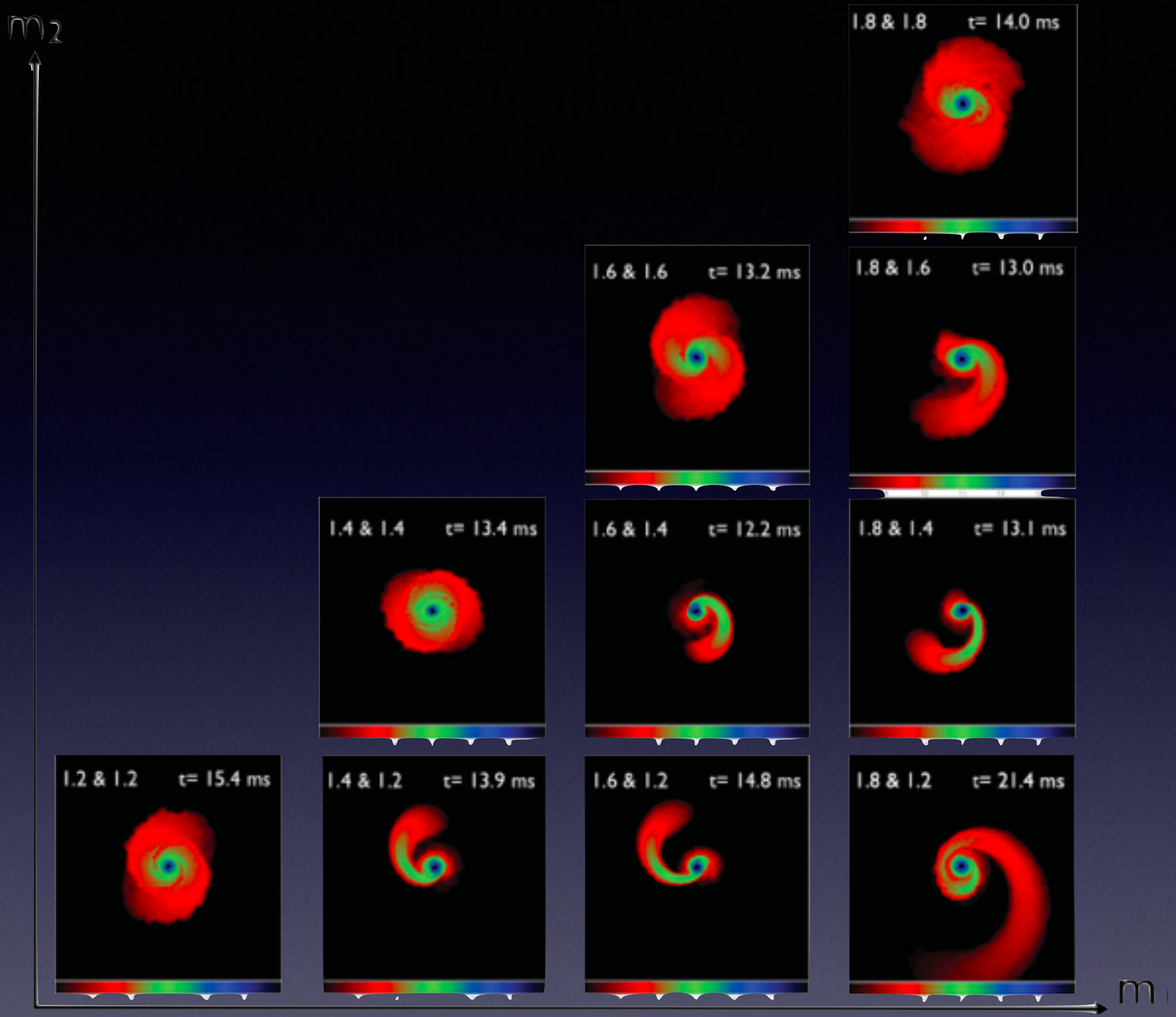
Photons escape from this region

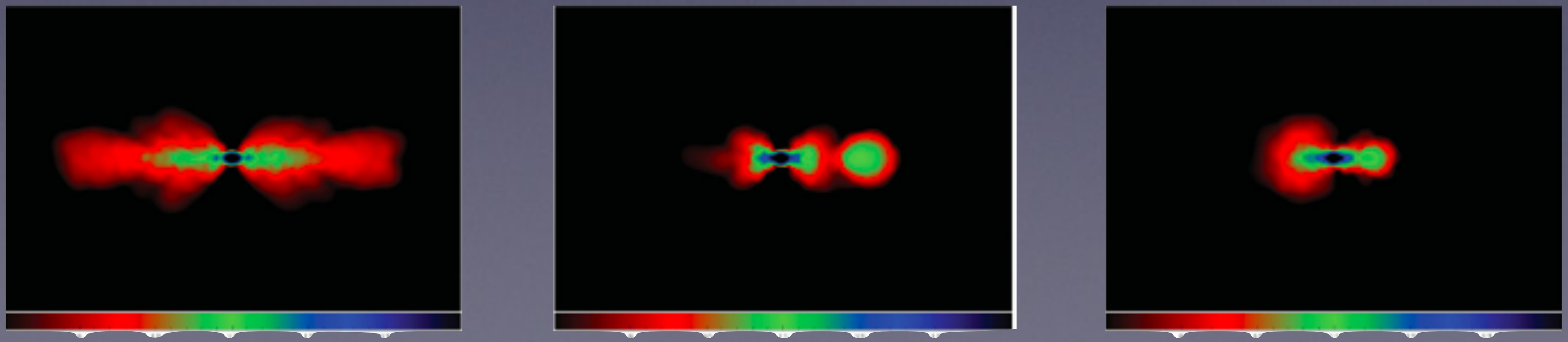


Increase as we see a large fraction of the matter.

Decrease due to radioactive decay

Macronova


m2
1.8 & 1.8 t= 14.0 ms
1.6 & 1.6 t= 13.2 ms
1.8 & 1.6 t= 13.0 ms
1.4 & 1.4 t= 13.4 ms
1.6 & 1.4 t= 12.2 ms
1.8 & 1.4 t= 13.1 ms
1.2 & 1.2 t= 15.4 ms
1.4 & 1.2 t= 13.9 ms
1.6 & 1.2 t= 14.8 ms
1.8 & 1.2 t= 21.4 ms
m1



Peak time and peak luminosity

Diffusion time = expansion time <=>
Mass of the “emitting region”

$$\frac{m(v)}{v} = \frac{4\pi c t^2}{\kappa}$$

Luminosity

$$L(t) = \dot{\epsilon}(t) m(v) = \dot{\epsilon}_0 (t/t_0)^{-\alpha} m(v)$$

Radioactive heating rate

The peak time

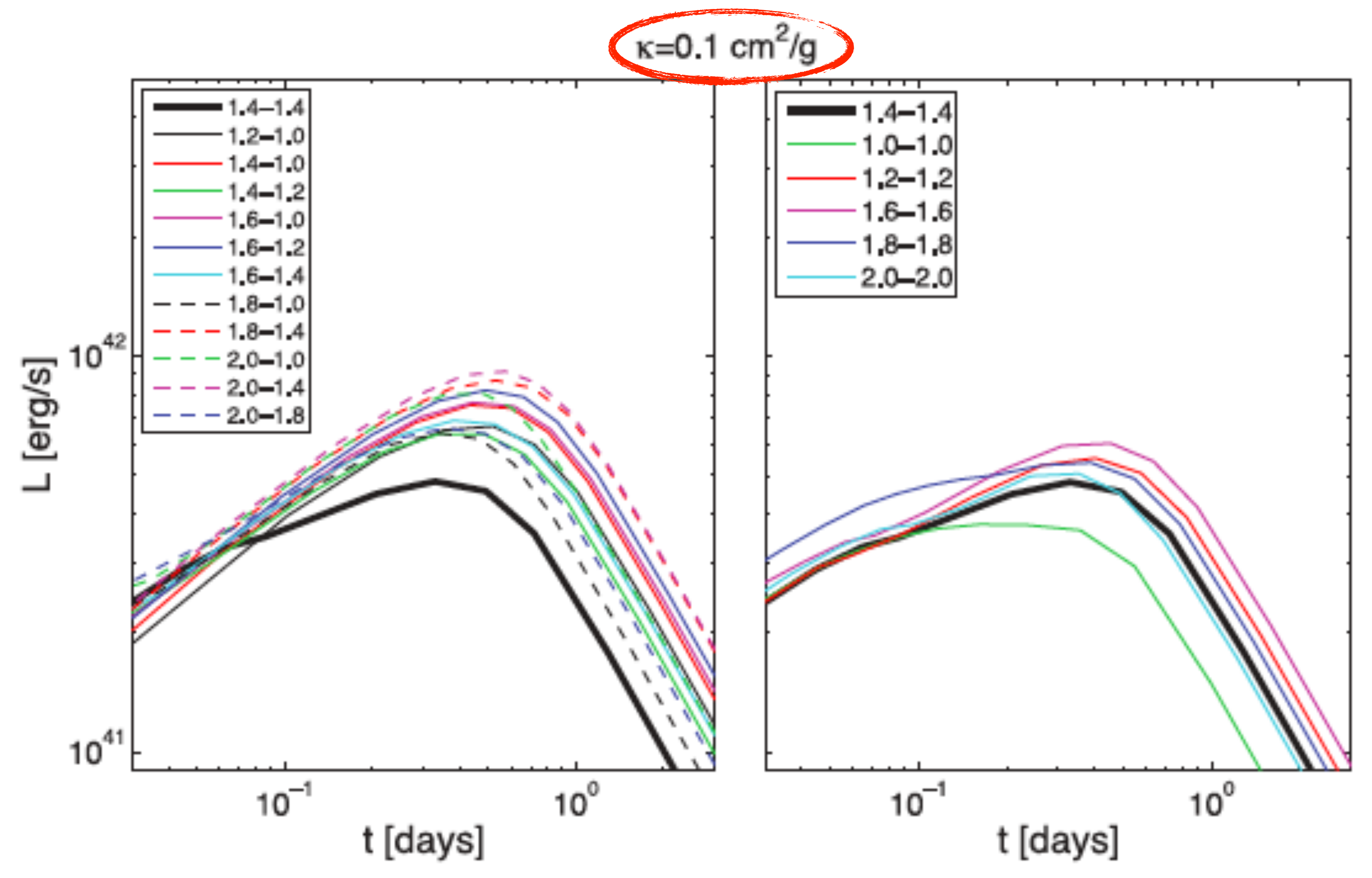
$$\tilde{t}_p \approx \sqrt{\frac{\kappa m_{\rm ej}}{4\pi c \bar{v}}} = 4.9 \text{ days} \left(\frac{\kappa_{10} m_{{\rm ej},-2}}{\bar{v}_{-1}}\right)^{1/2}$$

The peak luminosity

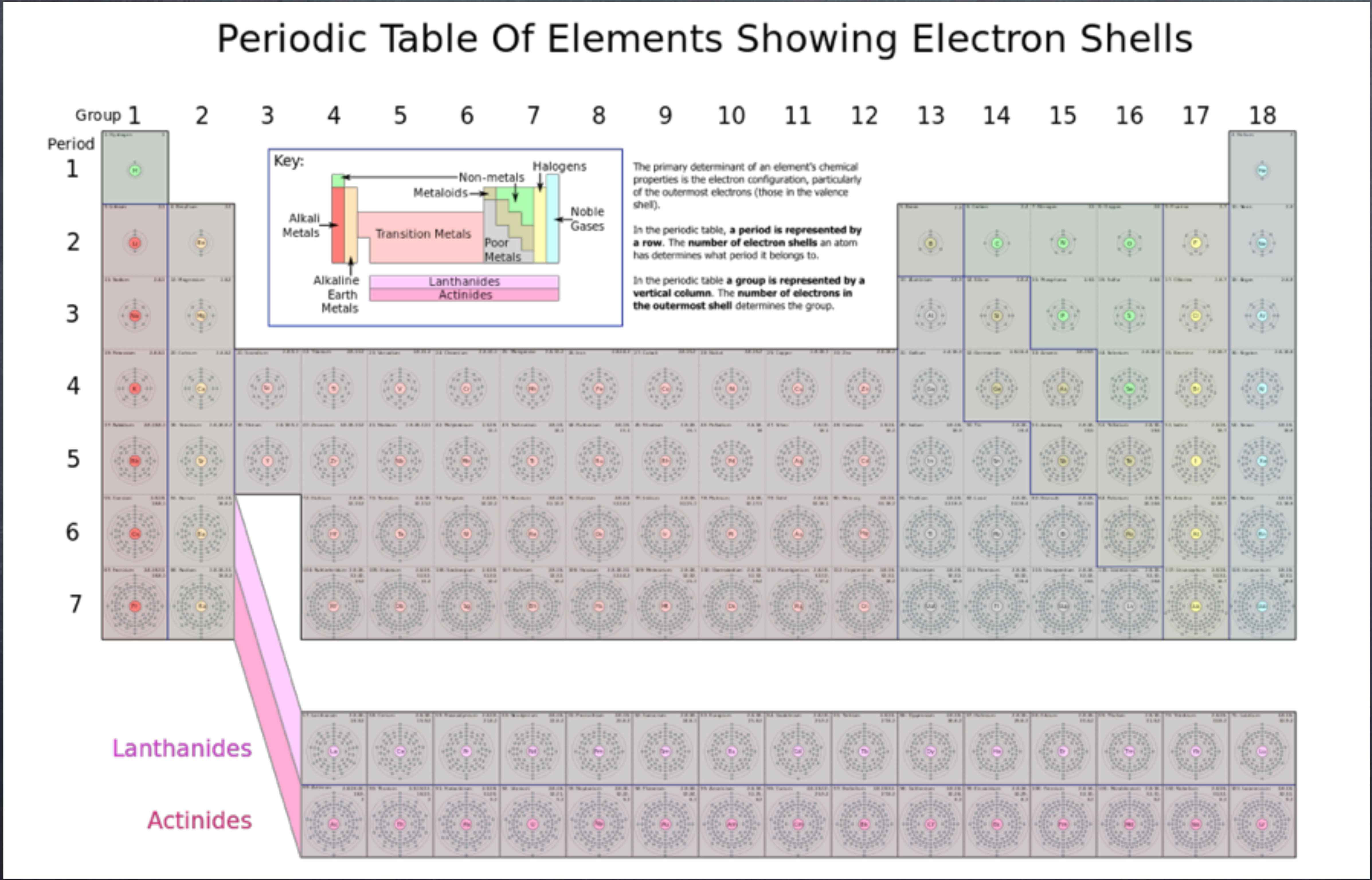
$$\tilde{L}_p \approx \dot{\epsilon}_0 m_{\rm ej} \left(\frac{\kappa m_{\rm ej}}{4\pi c \bar{v} t_0^2}\right)^{-\alpha/2} = 2.5 \times 10^{40} \frac{\rm erg}{\rm s} \left(\frac{\bar{v}_{-1}}{\kappa_{10}}\right)^{\alpha/2} m_{{\rm ej},-2}^{1-\alpha/2}$$

Macronova light curves

Metzger et al., 2011; TP, Nakar, Rosswog, 13



Lanthanides



Periodic Table Of Elements Showing Electron Shells

The primary determinant of an element's chemical properties is the electron configuration, particularly of the outermost electrons (those in the valence shell).

In the periodic table, **a period is represented by a row**. The **number of electron shells** an atom has determines what period it belongs to.

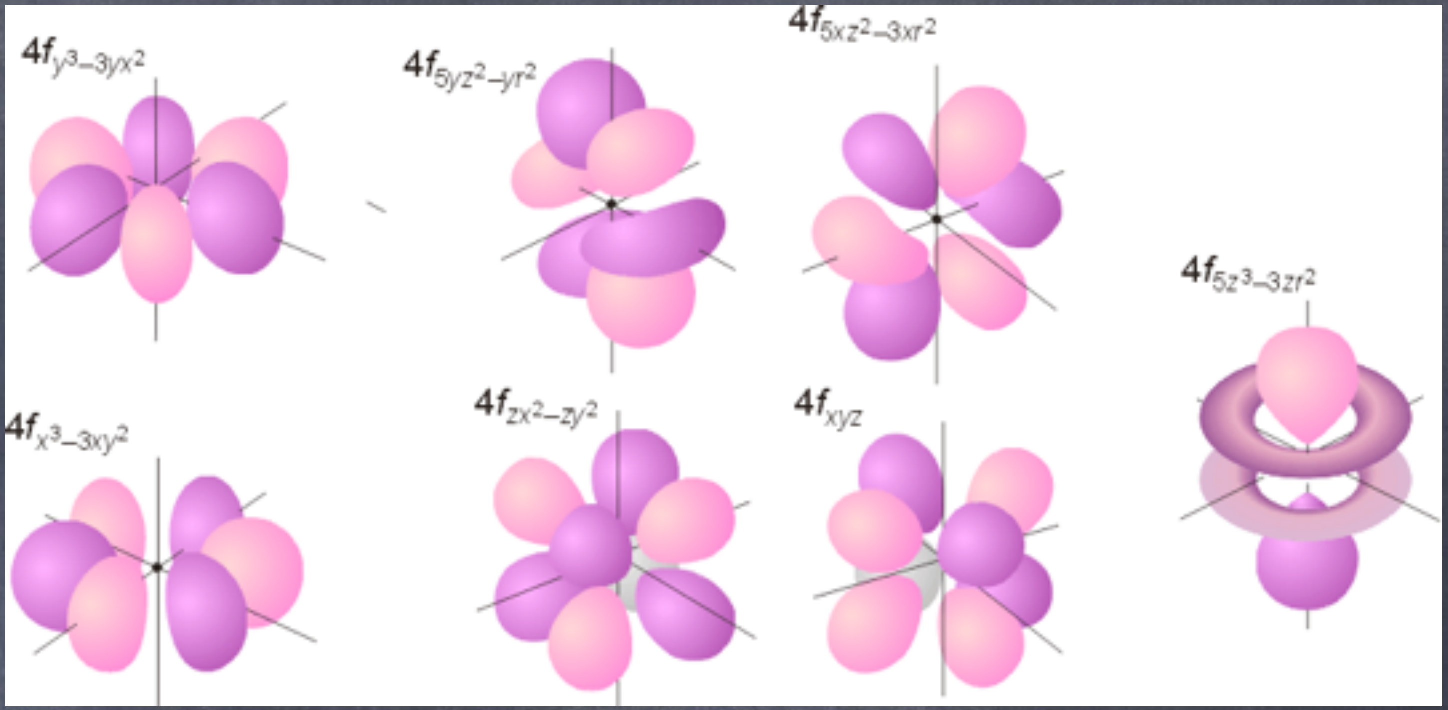
In the periodic table **a group is represented by a vertical column**. The **number of electrons in the outermost shell** determines the group.

Why do are the Lanthanides “out” of the table?

The Lanthanides' Opacity

Kassen & Barnes 2013

The Lanthanides have "too many" lines

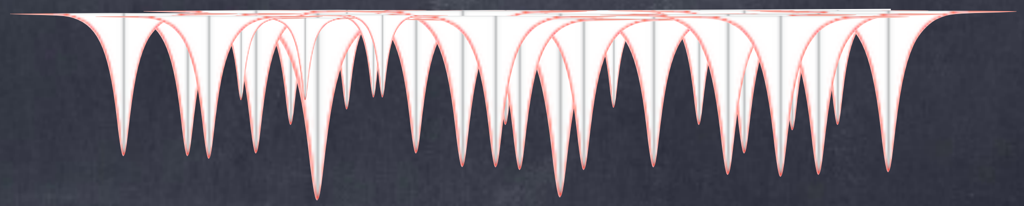


κ=10 cm^2/gm

compare with

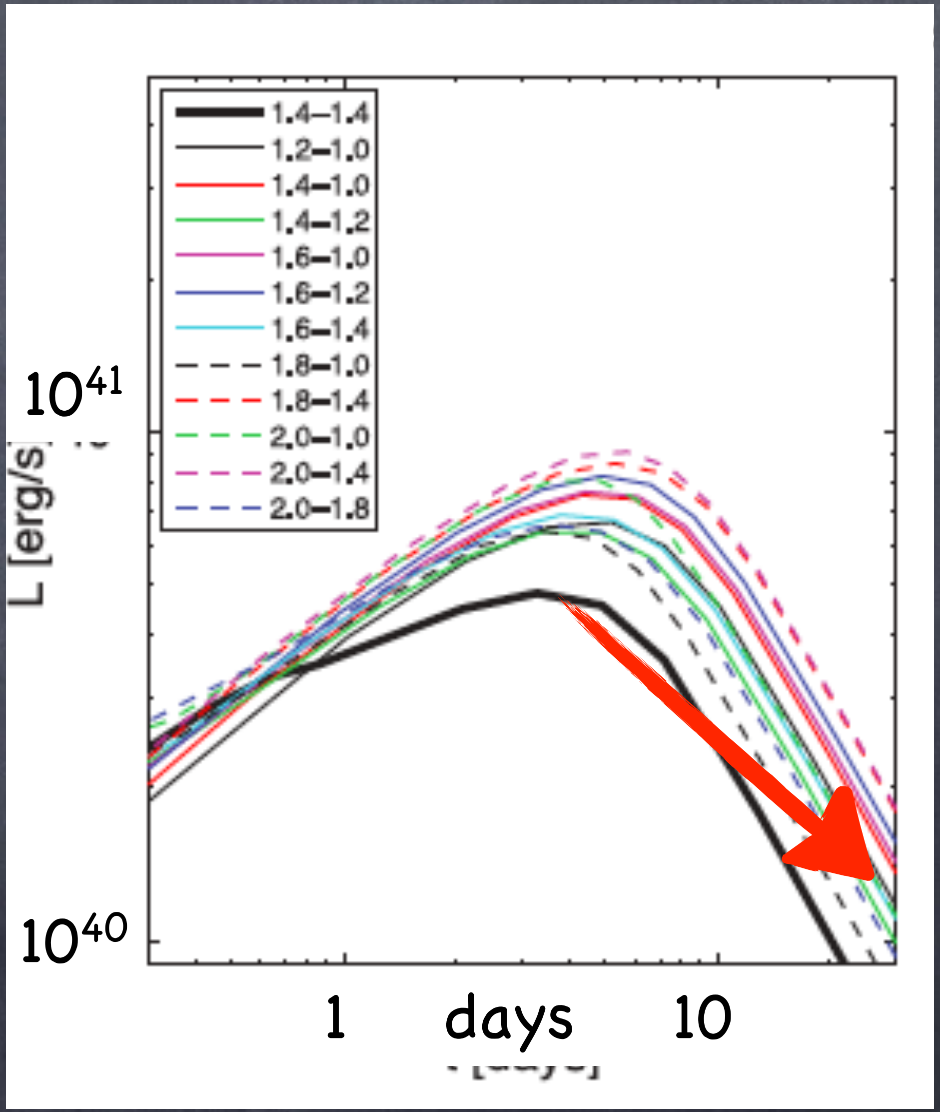
κ =0.4 cm^2/gm for the iron group

κ_T=0.1 cm^2/gm for electron scattering



Lanthanides dominate the Opacity (Kassen & Barnes 13)

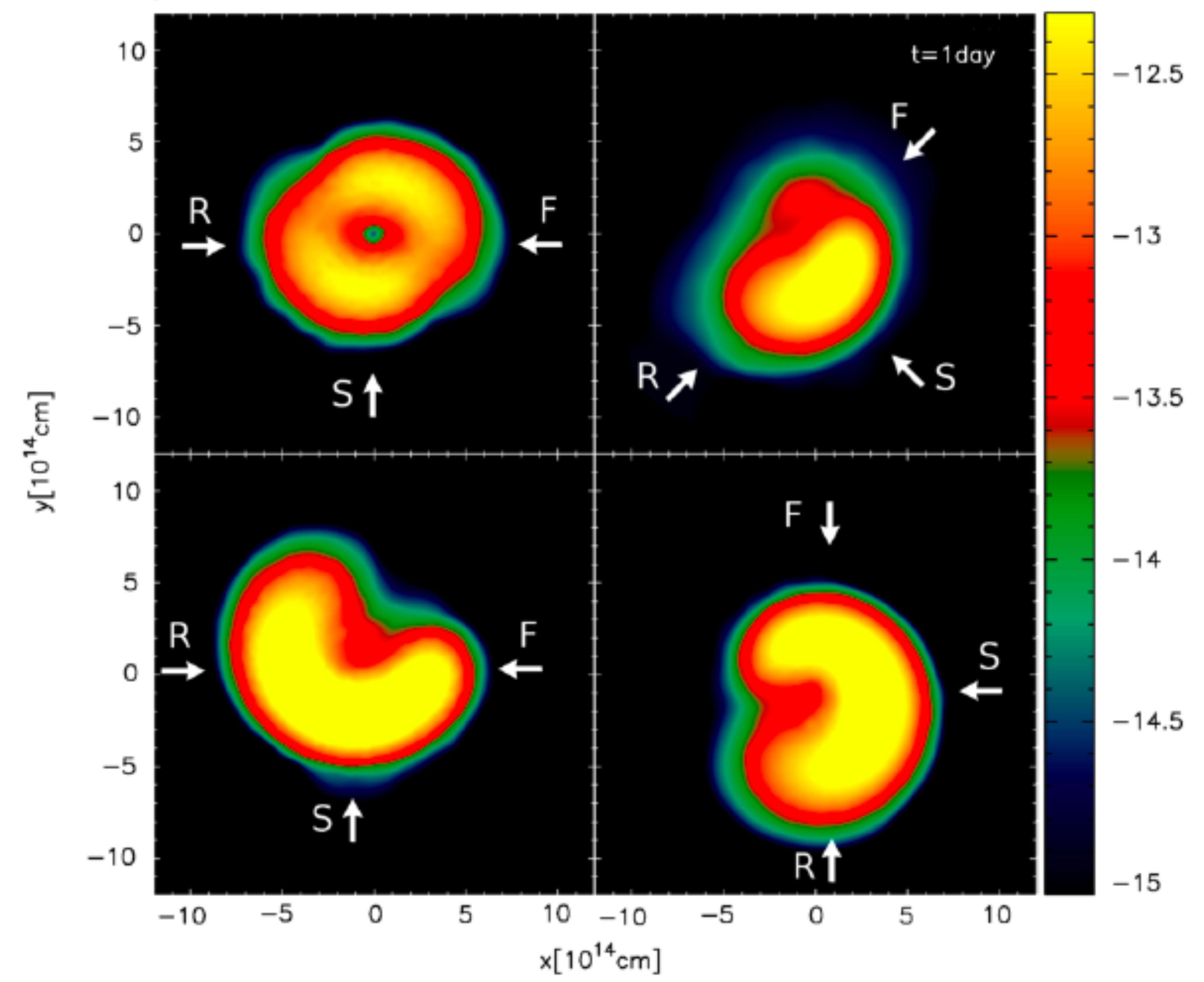
- $\kappa = 10\,\mathrm{cm}^2/\mathrm{gm}$
- $t_{max} \propto \kappa^{1/2}$ => longer
- $L_{max} \propto \kappa^{-0.65}$ => weaker
- $T \propto \kappa^{-0.4}$ => redder



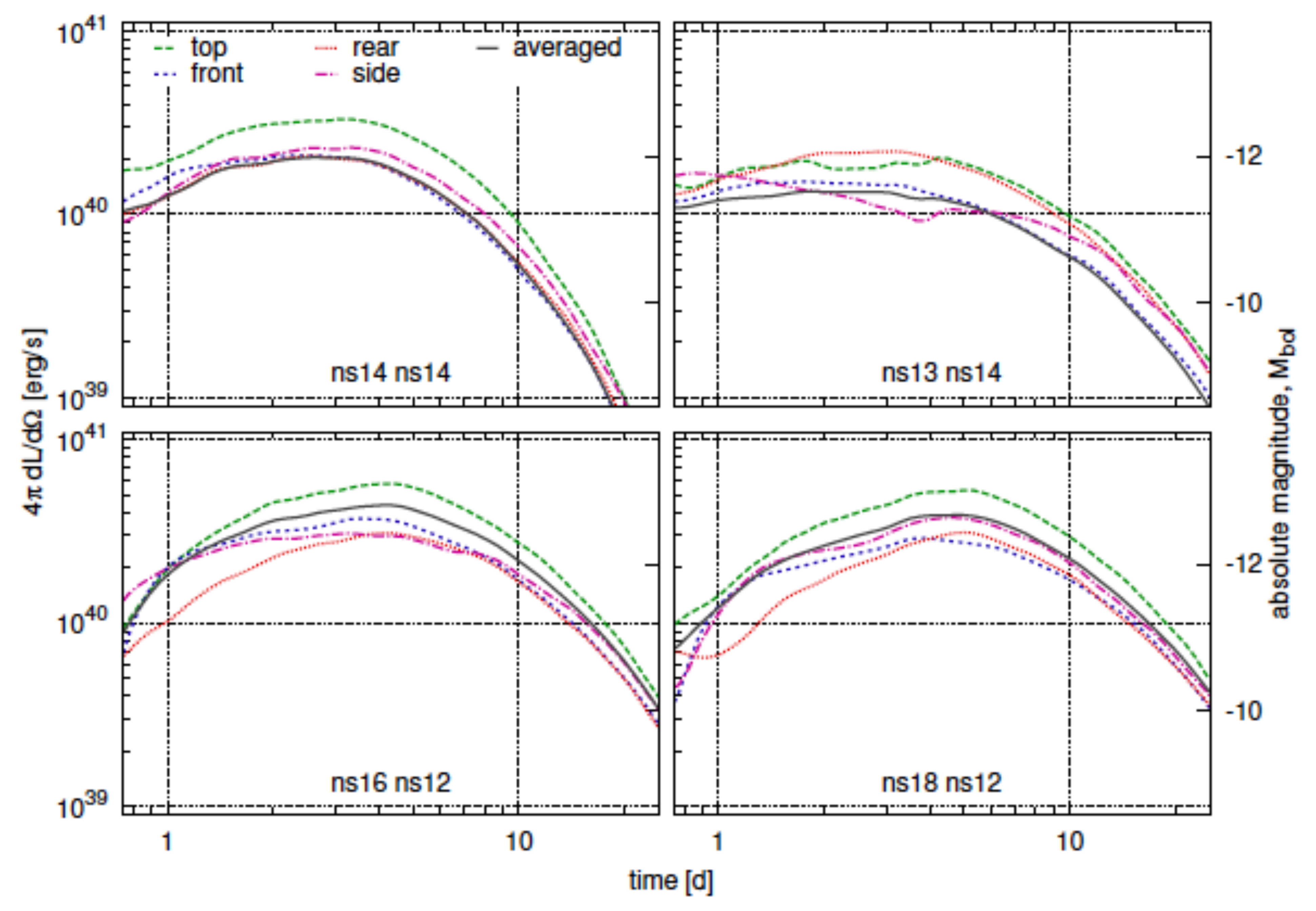
uv or optical -> IR

More detailed estimates

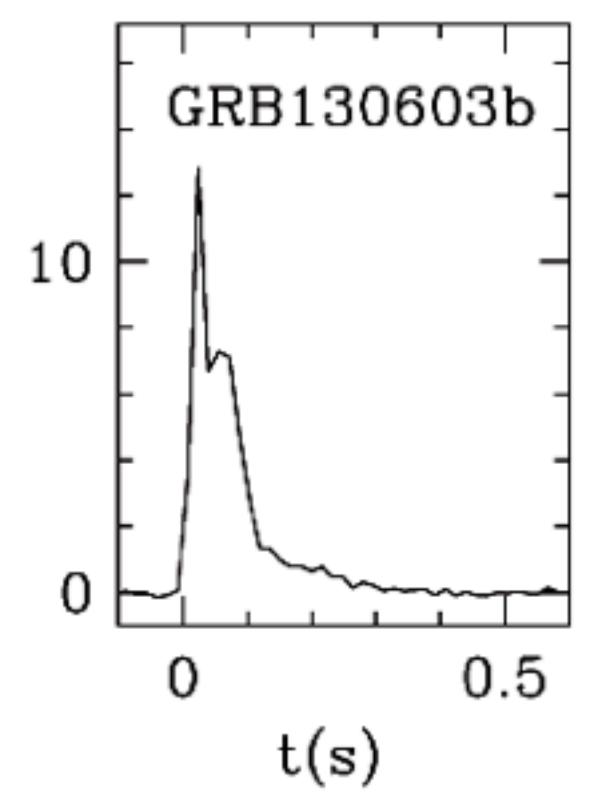
Grossman, Korobkin TP Rosswog, 13

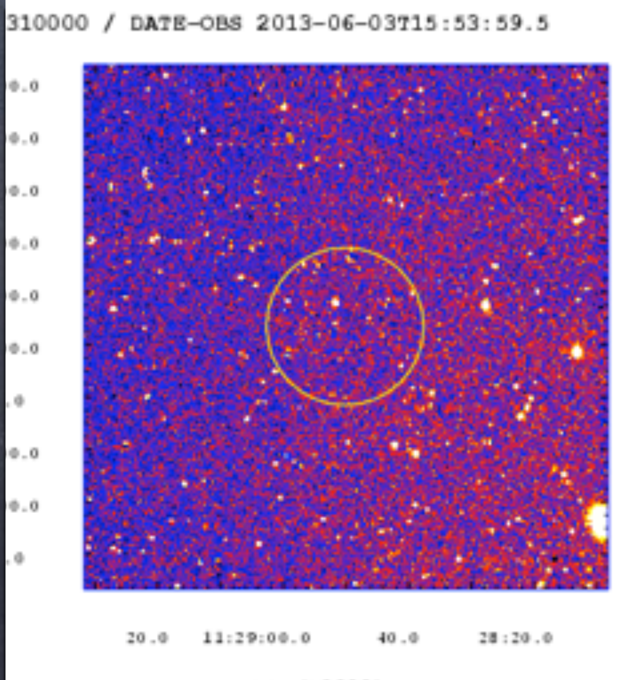


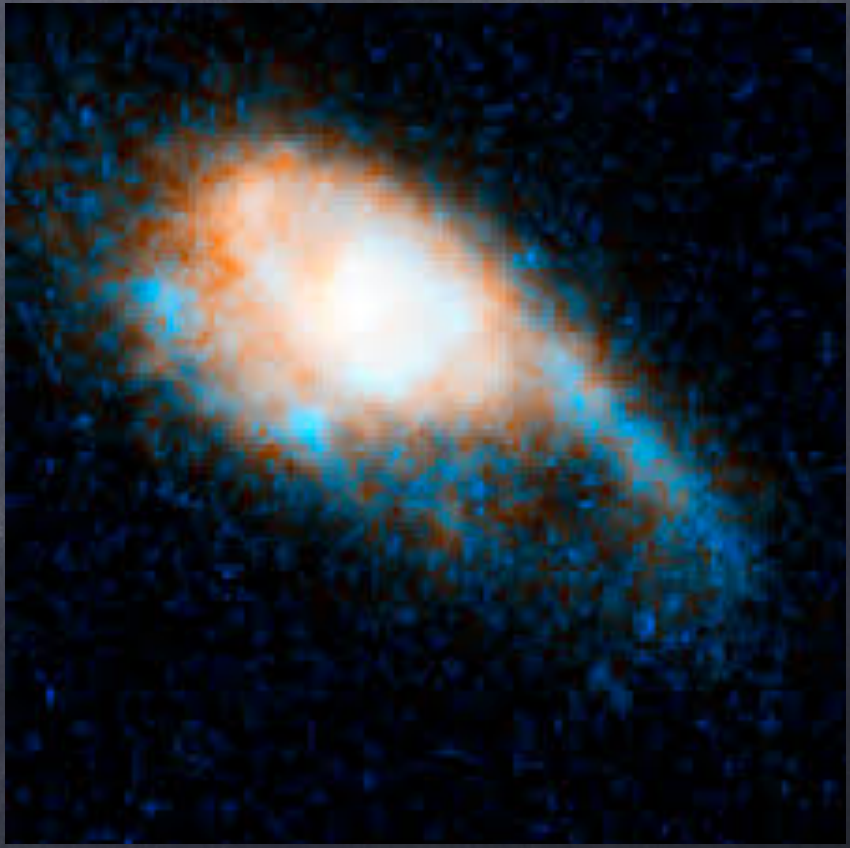
Bolometric light curves



Putting it all together
5. Gamma-Ray Burst (GRB)
130603B

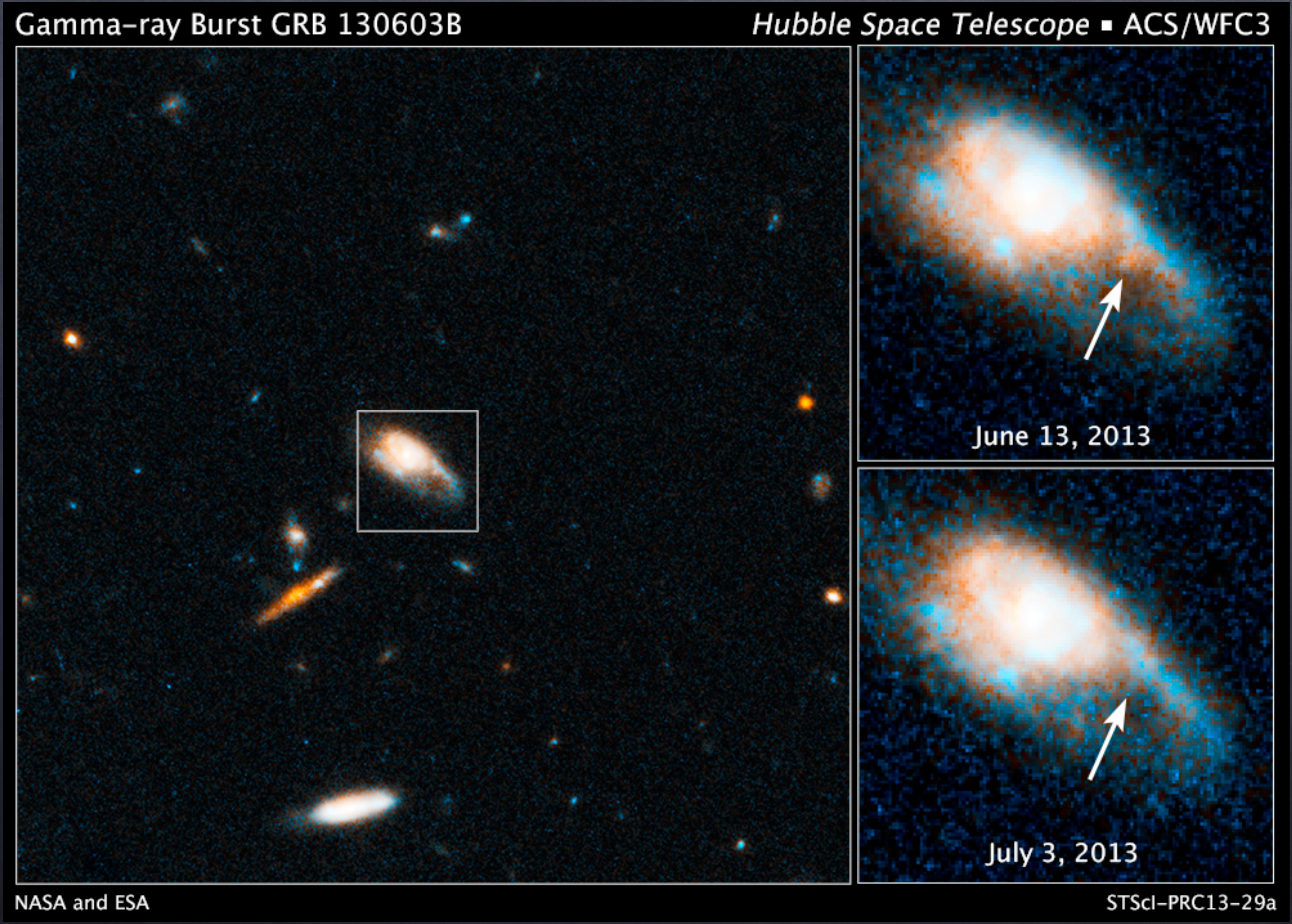






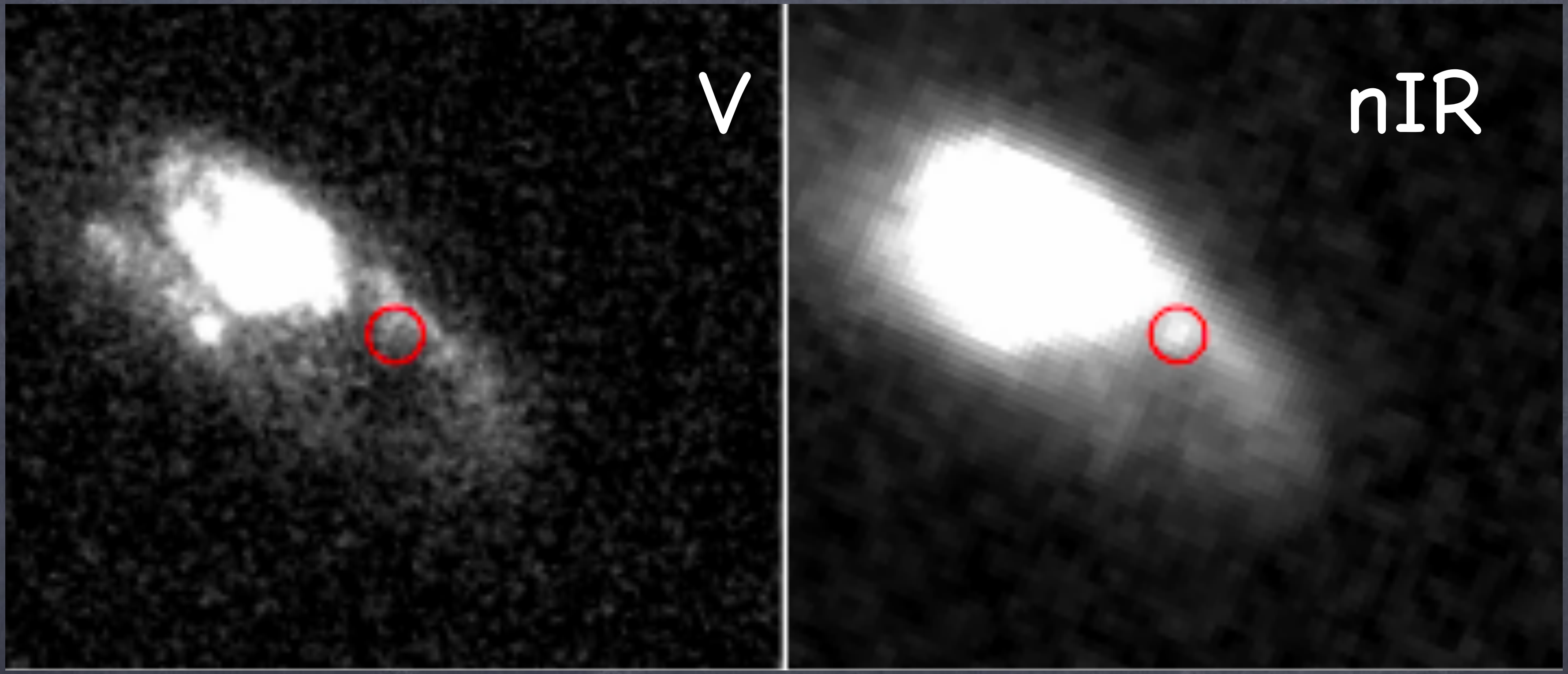
GRB 130603B

z=0.356 <=> 1 Gpc = 3 Glyr

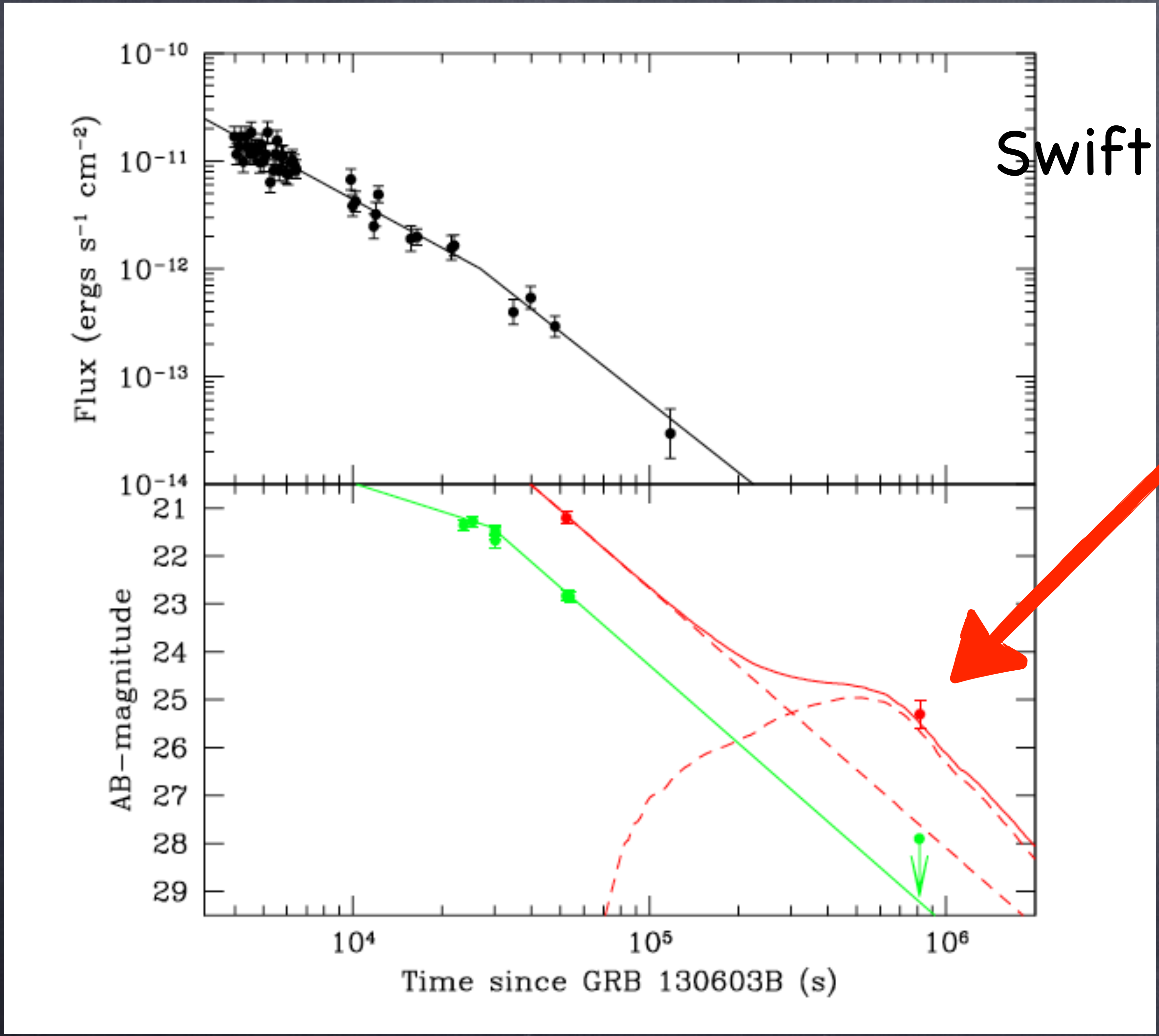

Gamma-ray Burst GRB 130603B
Hubble Space Telescope ▪ ACS/WFC3
June 13, 2013
July 3, 2013
NASA and ESA
STScI-PRC13-29a

GRB130603B @ 9 days AB

(6.6 days at the source frame)



HST image (Tanvir + 13)



Swift

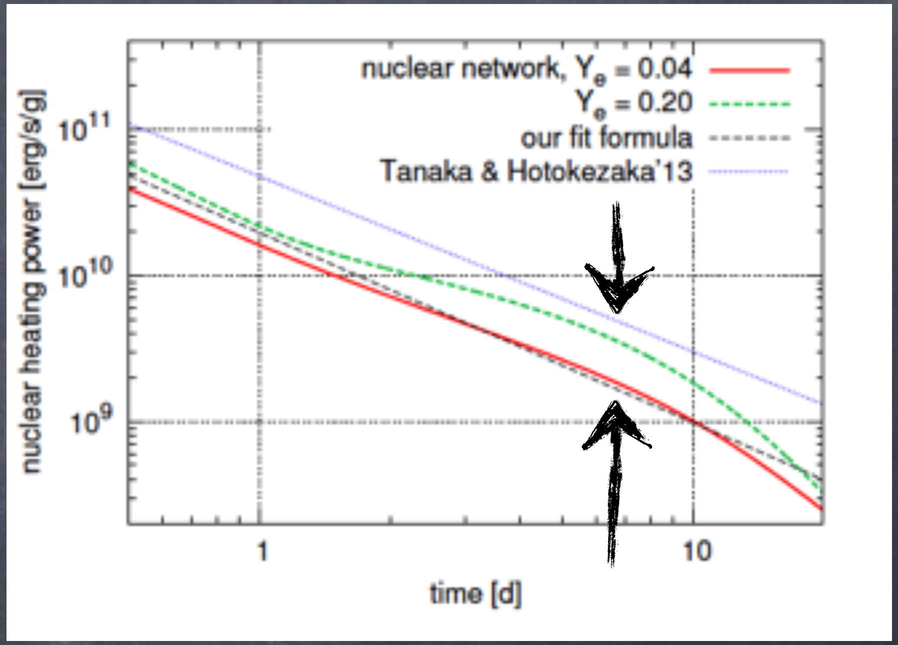
Macronova?

Tanvir + 13

GRB130603B @ z=0.356 nIR transient

- Consistent with Barnes & Kasen (13) and Tanaka & Hotozoka (13)

- But Both groups possibly overestimated radioactive heating rate by a factor of 2-4

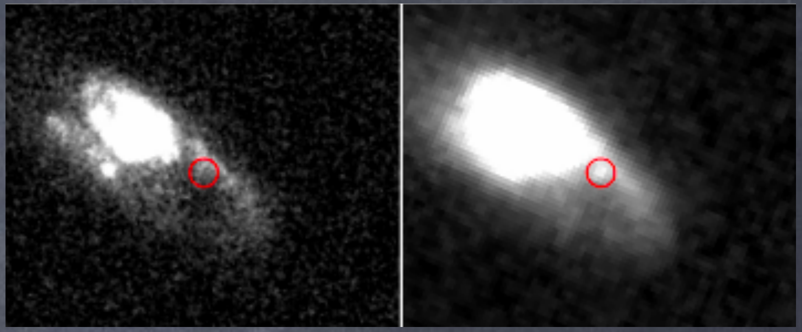


- The expected signal is slightly too large

If correct

- Confirmaiton of the GRB neutron star merger model (Eichler, Livio, TP & Schramm 1989).
- Confirmation of the Li-Paczynski Macronova.
- Confirmation that compact binary mergers are the source of heavy (A>130) r-process material (Gold, Silver, Platinum, Plotonium, Uranium etc...).

6. The Origin of GOLD



Implications

Mass ejected in a merger

Observed luminosity = 10^{41}erg/sec @ 6.6 days

$$m_{ej} > 0.02(\epsilon/0.5)^{-1} m_{\odot}$$

# of mergers

$$N = 2.5 \times 10^5 \left(\frac{M^{A>130}}{10^4 m_{\odot}}\right)\left(\frac{m_{ej}}{0.04 m_{\odot}}\right)^{-1}$$

A>130 r-process material in the Galaxy

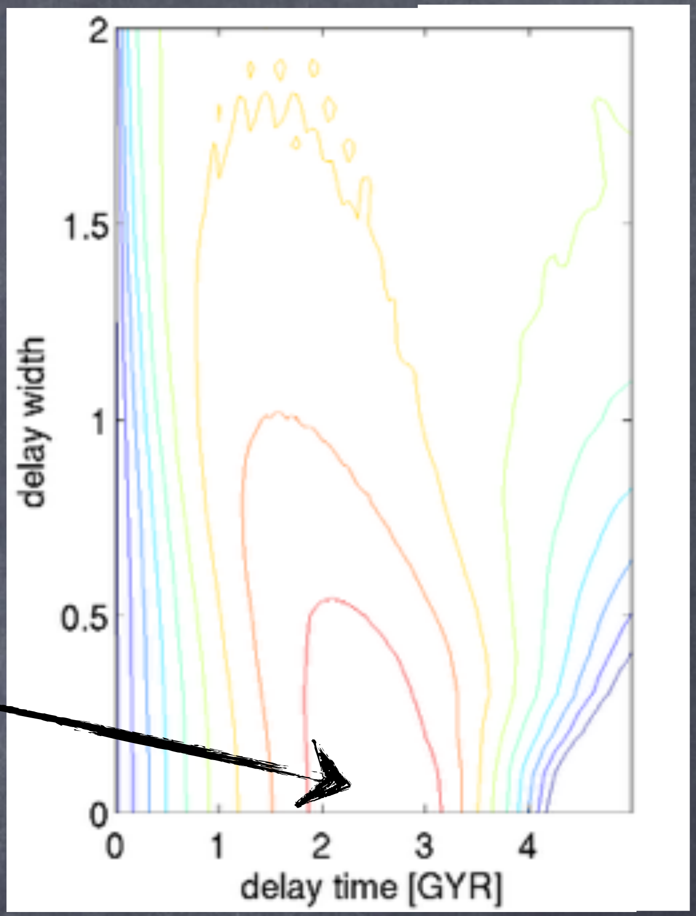
Mergers' Rate

$$\begin{aligned} R_{merger} &= 20 \left(\frac{m_{ej}}{0.04 m_{\odot}}\right)^{-1}\left(\frac{M^{A>130}}{10^4 m_{\odot}}\right) \mathrm{Myr}^{-1} \\ &= 200 \left(\frac{m_{ej}}{0.04 m_{\odot}}\right)^{-1}\left(\frac{M^{A>130}}{10^4 m_{\odot}}\right) \mathrm{Gpc}^{-3}\mathrm{yr}^{-1} \end{aligned}$$

The rate of short GRBs

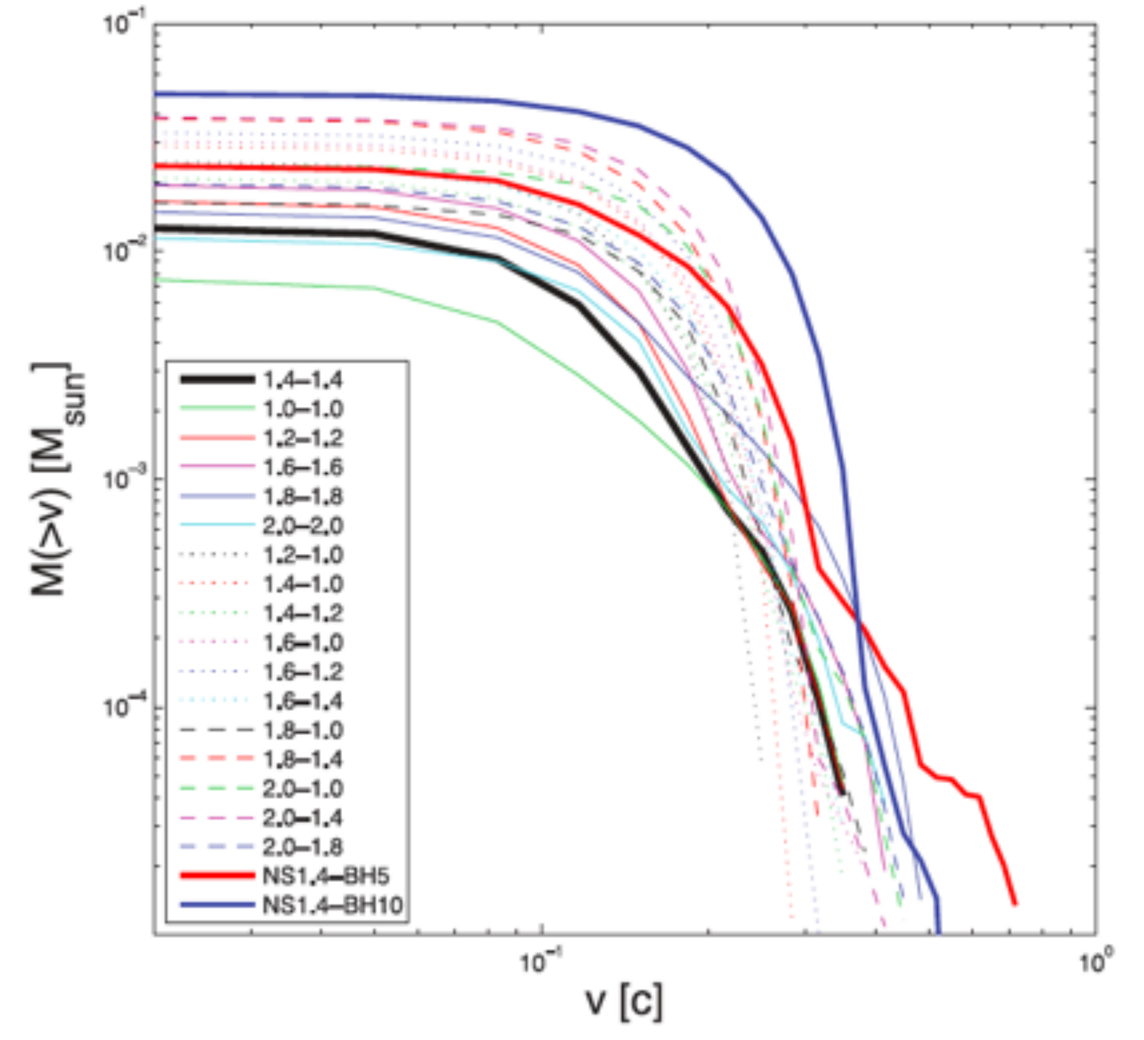
Guetta & TP 2006; Wanderman & TP 2013

- $R_{sgrb} = 5 \pm 2\ \mathrm{Gpc}^{-3}\ \mathrm{yr}^{-1}$
- Typical spiral-in phase of 2.5 Gyr.
- Consistent with $R_{merger} = 200\ \mathrm{Gpc}^{-3}\ \mathrm{yr}^{-1}$ for a reasonable beaming factor of 40.
- Consistent with rate estimates based on galactic neutron star binaries.

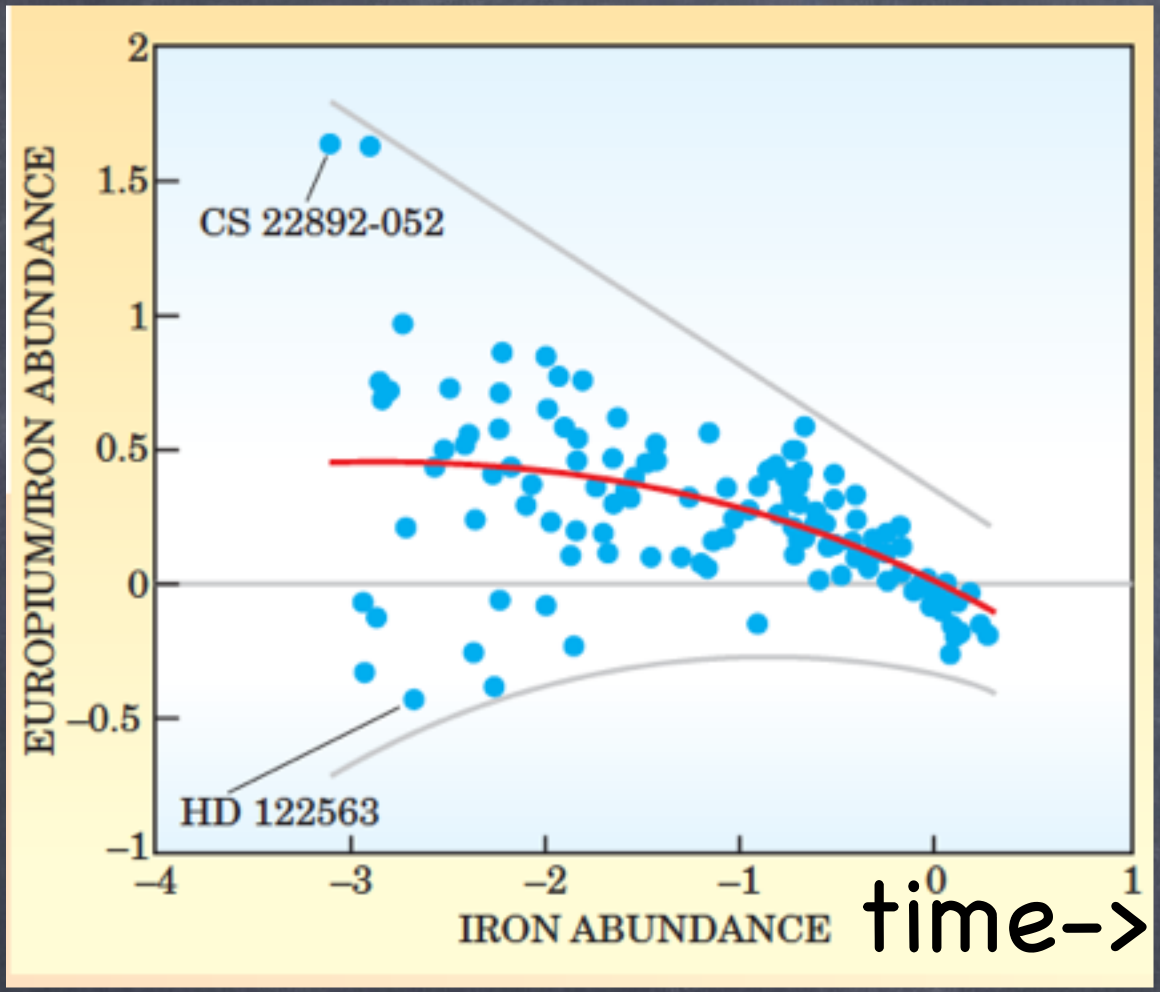


But:

- The ejected mass is about 0.04 M_{sun}. The minimal mass is 0.02 M_{sun}.
- This is rather large for neutron star binary merger.
- Is the solution black hole - neutron star merger?



Early nucleosynthesis - a challenge



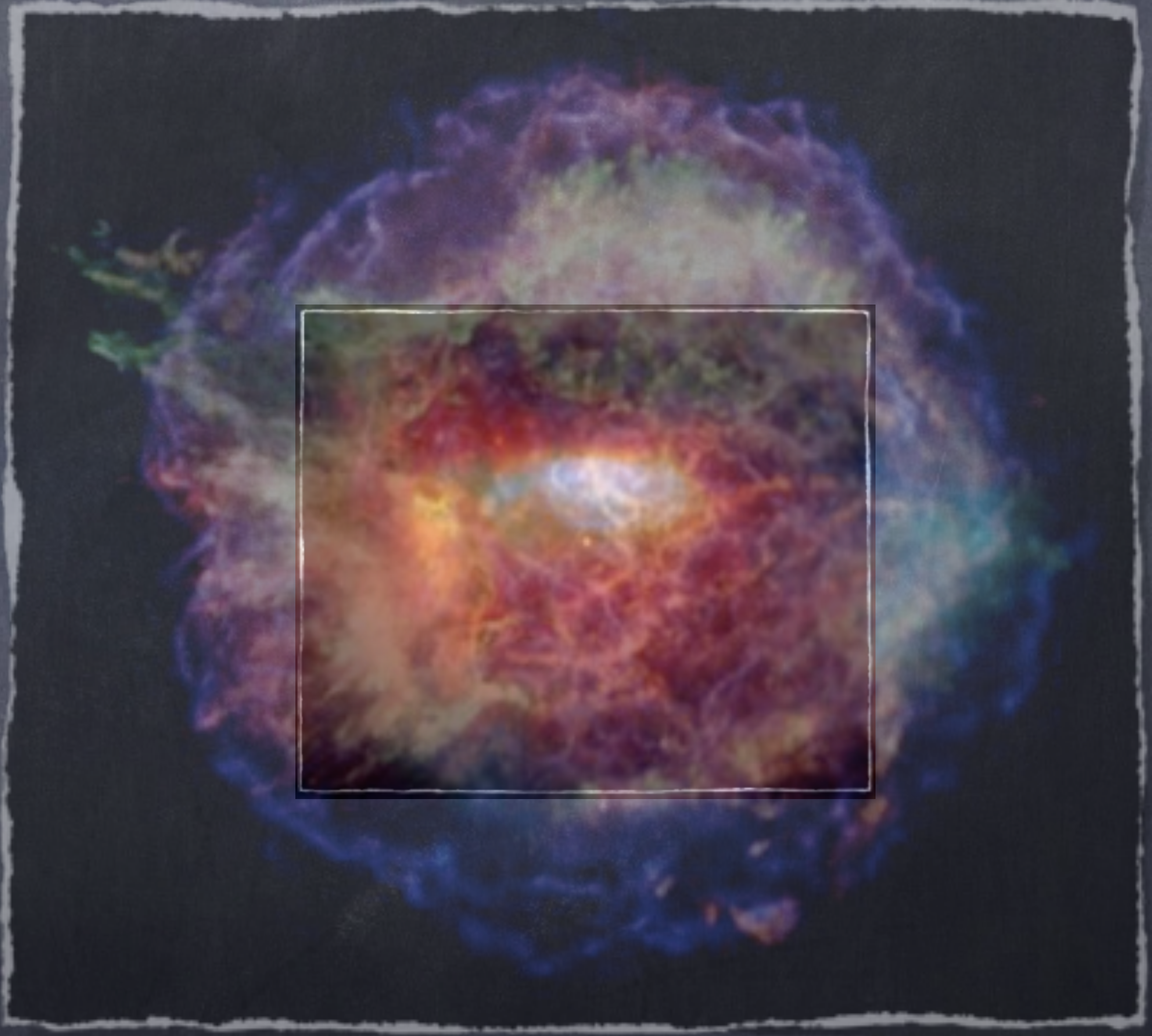
A population of fast mergers?

Figure 6. Europium abundance in a large sample of old and young stars, age being inferred from Fe abundance. The halo star HD 122563 is almost as Fe-poor as CS 22892-052, and therefore presumably just about as old, but it has much less Eu, an element made only in the r-process. The red line is a least-square-fit to the data, and the gray flanking curves indicate decreasing scatter in the data with increasing time. Numerical conventions are as in figure 5. Zero on the abscissa means Fe abundance like that of the 4.6-billion-year-old Sun.

From Cowan and Thielemann

The radio - flare (Nakar & Piran 2011)
Testing the Macronova interpretation

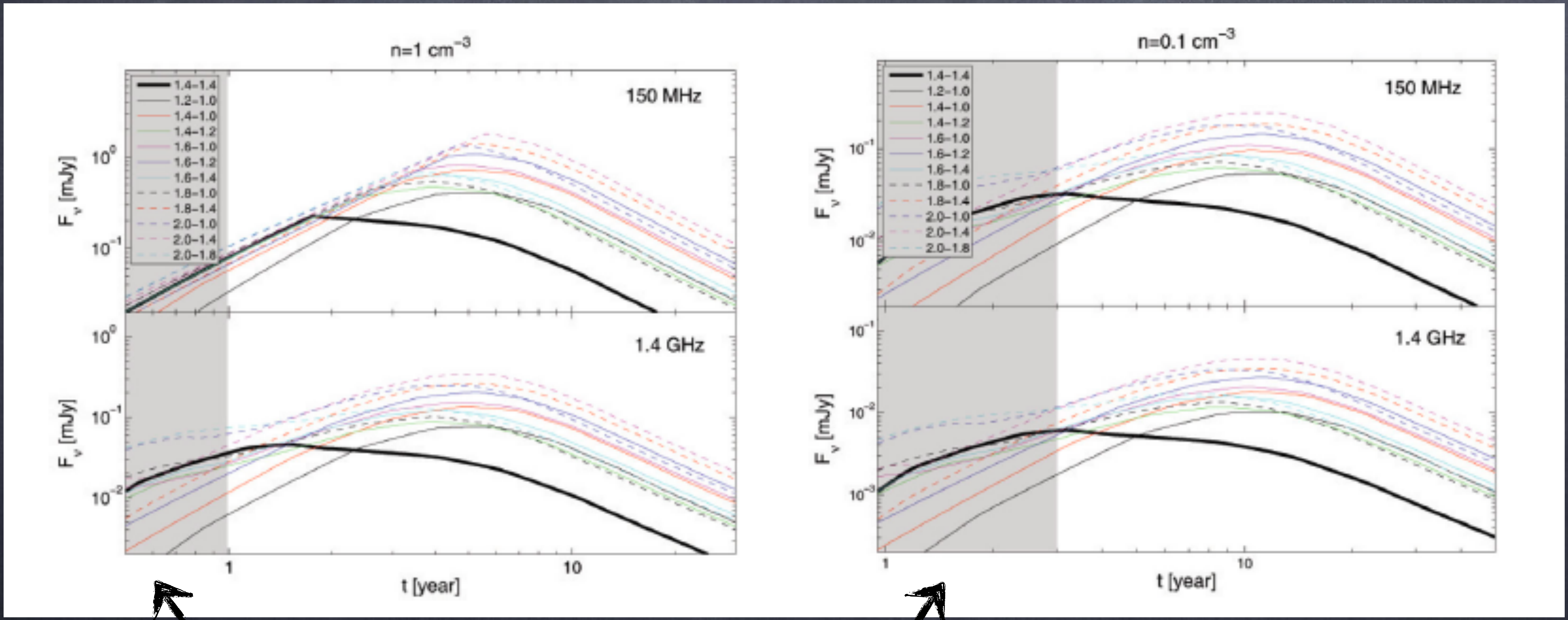
A long lasting radio flare due to the interaction of the ejecta with surrounding matter may follow the macronova.



Supernova -> Supernova remnant

Macronova -> Radio Flare

Radio frlares from neutron star mergers

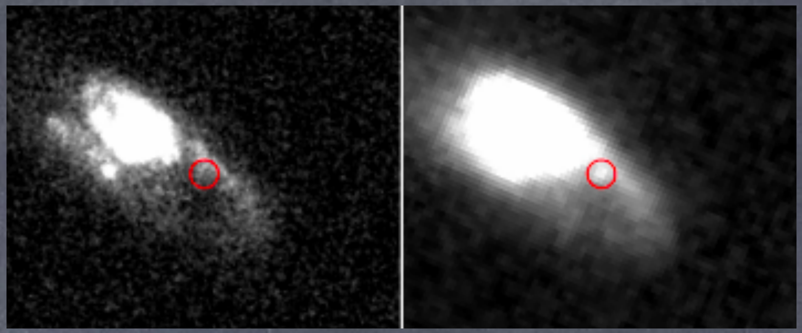


dominated by high velocity ejecta

A flare from GRB 130603B should be easily detected by the EVLA (if external density is not too small)

Summary

- **There are a few caveats -** But
- The nIR flare that followed the short GRB 130603B could have been a Macronova. If so than:
  - ✓Short GRBs arise from mergers.
  - ✓Gold and other A>130 elemets are produced in mergers. (But large m_{ej} and short time delay).
- A radio flare may confirm this!
- Another strong well localized short GRB is expected within a year or so.



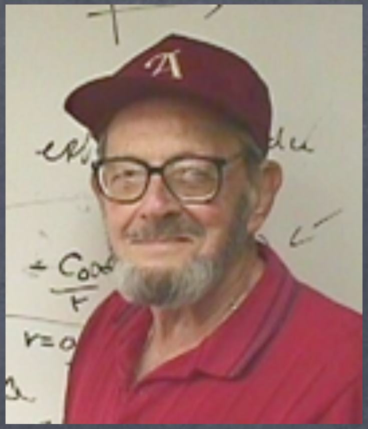
And ->

One cannot give a talk in Astronomy these days without a reference to the Solar System and life.

- The early Solar System had ^{244}Pu (τ = 117 Myr)

  Wasserburg et al, (2006).

No evidence for ^{244}Pu deposition in deep-sea crust and sediment accumulated over the last ~25 Myr (M. Paul et al., 2001; A. Wallner et al., in preparation).
=> ^{244}Pu is NOT from the Inter Stellar Medium!
=> Actinides production near the early Solar System just prior to formation.



Gerry Wasserburg

- Irregular production from rare episodes.
  => E.g. a merger within <50 pc=150 lyr from the solar system just prior to its formation?

The End ?